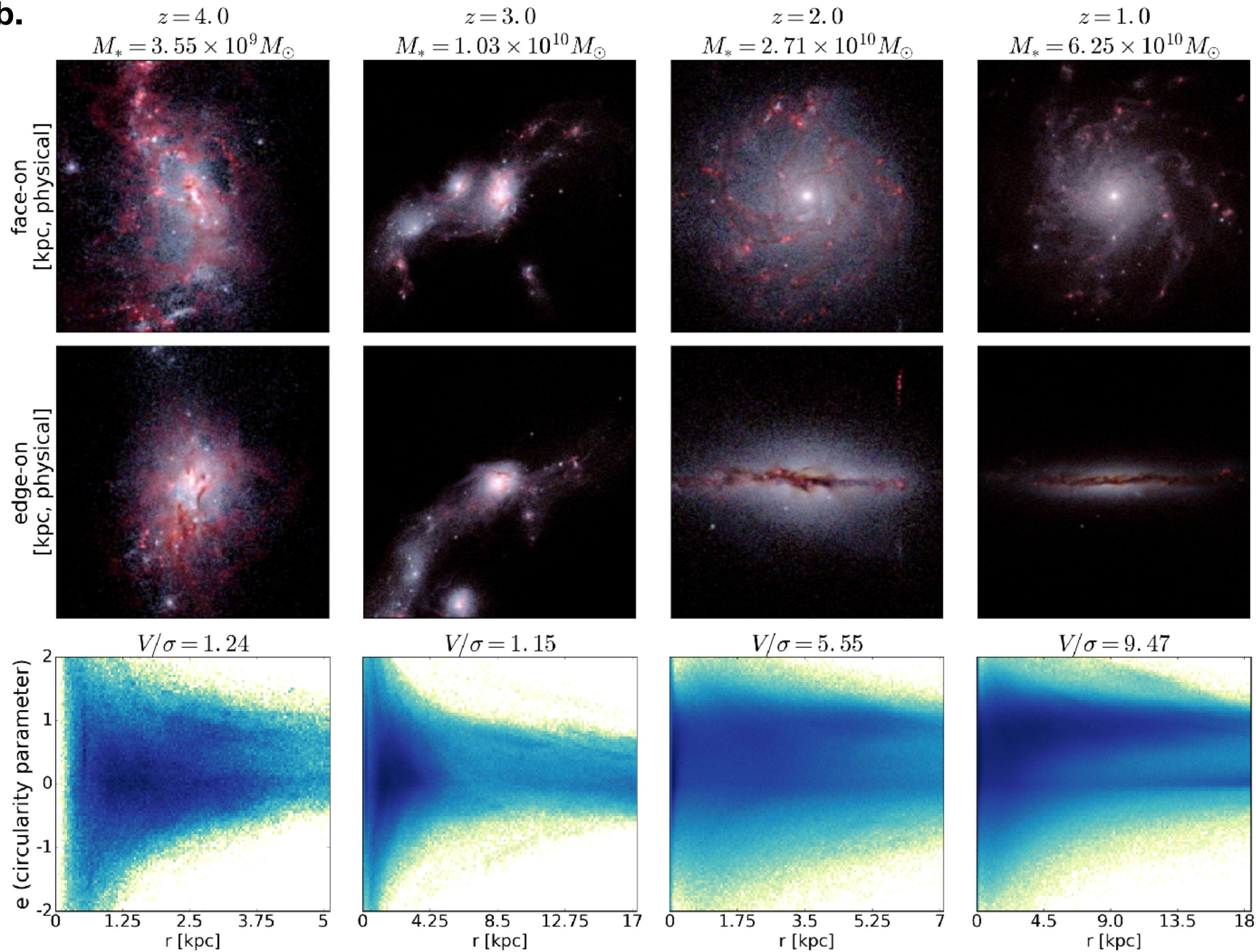


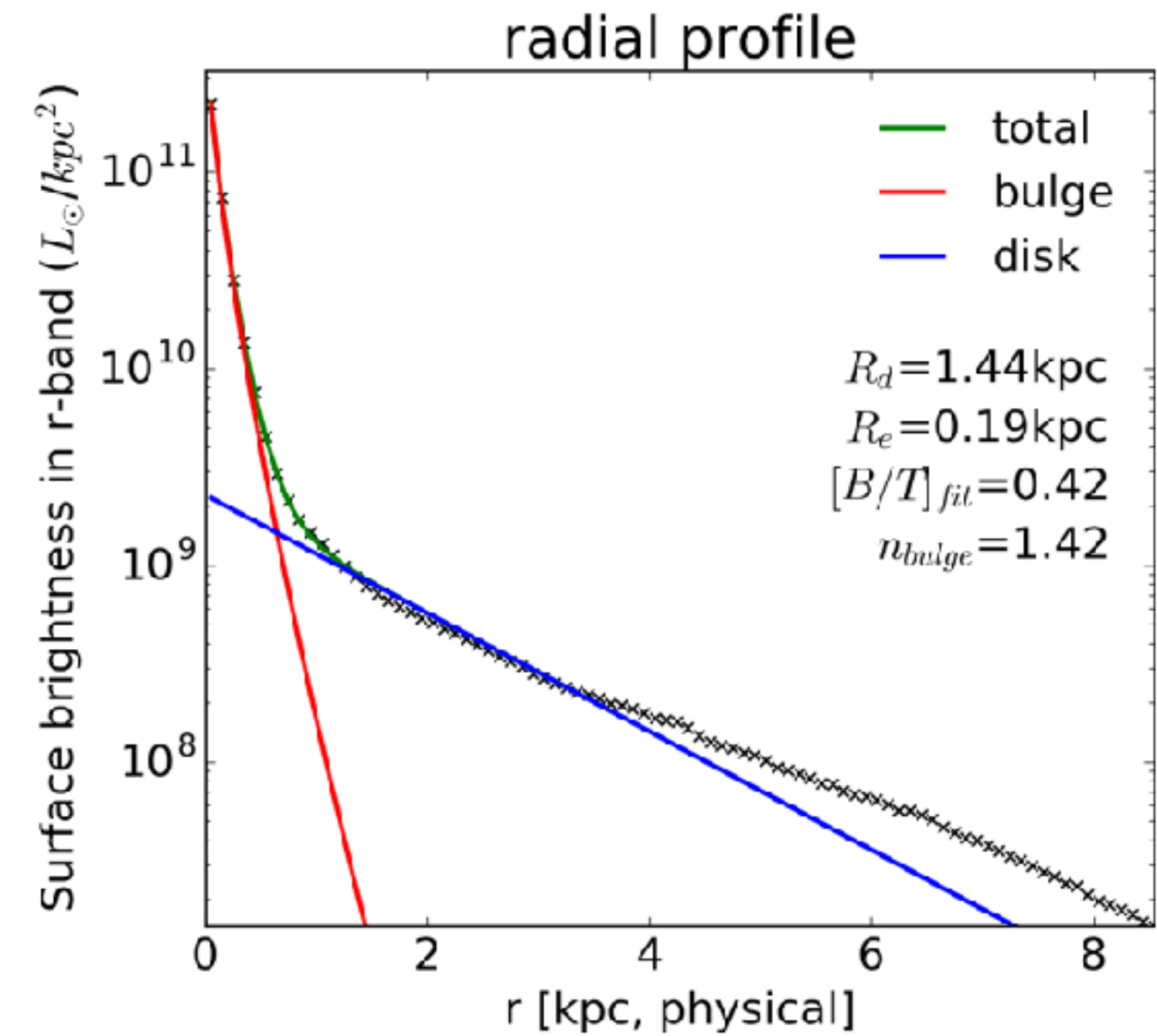
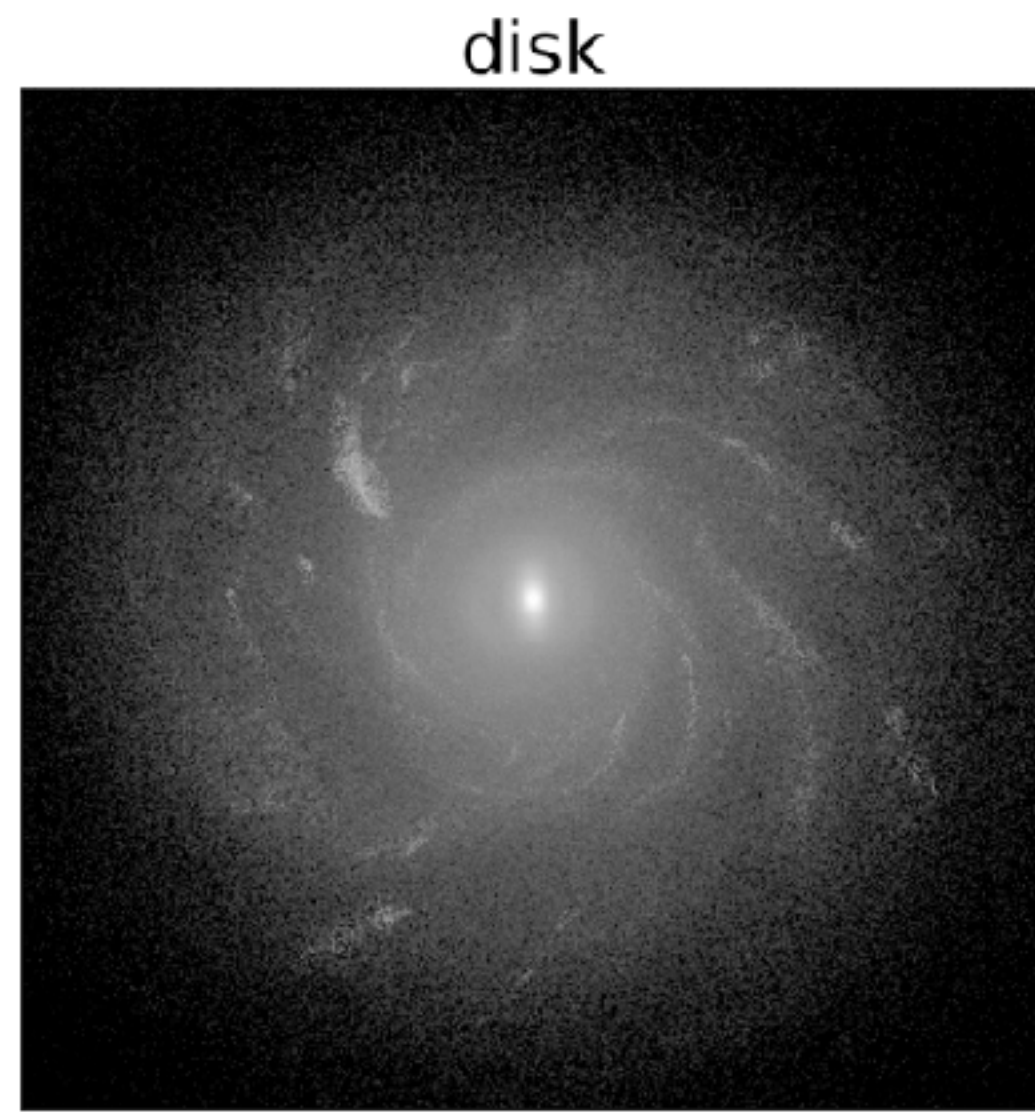
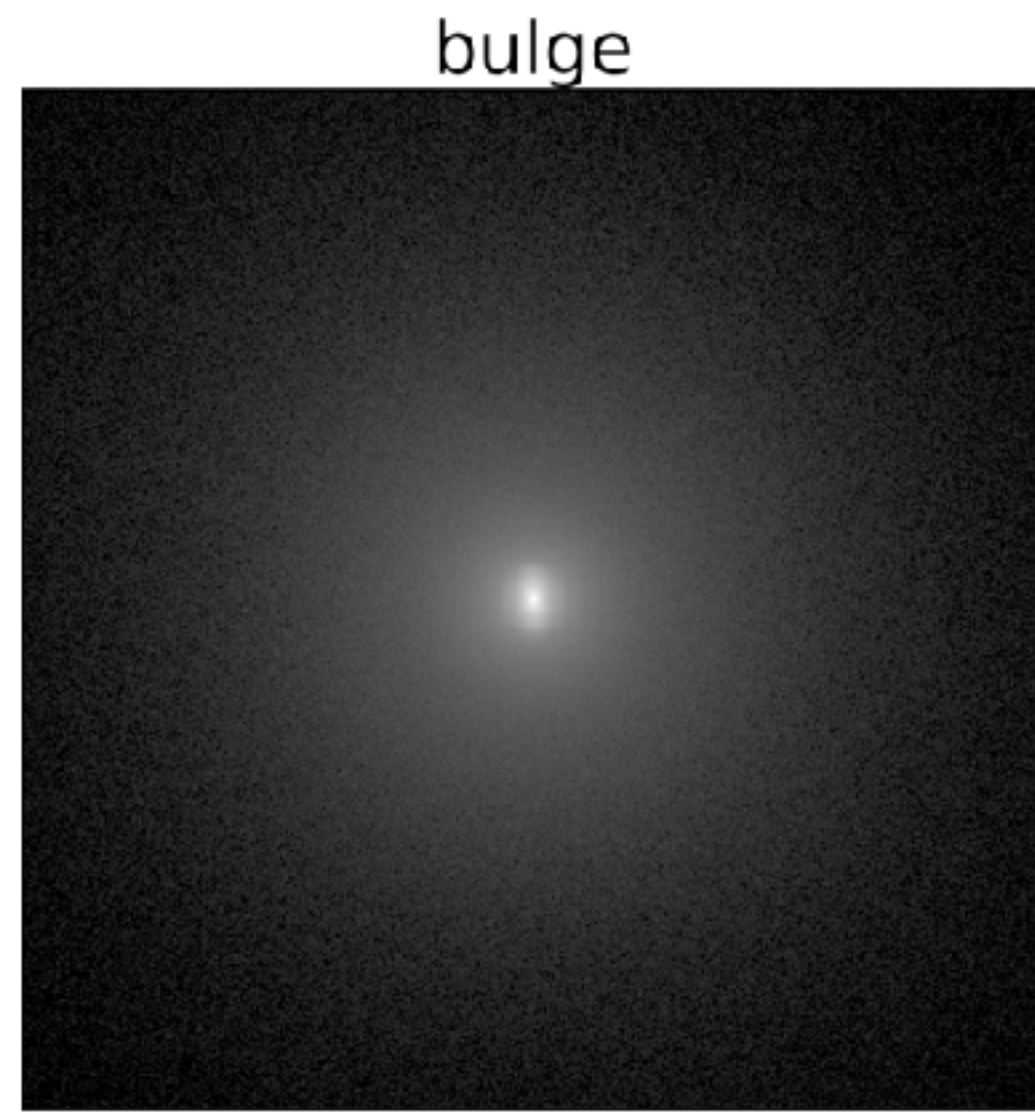
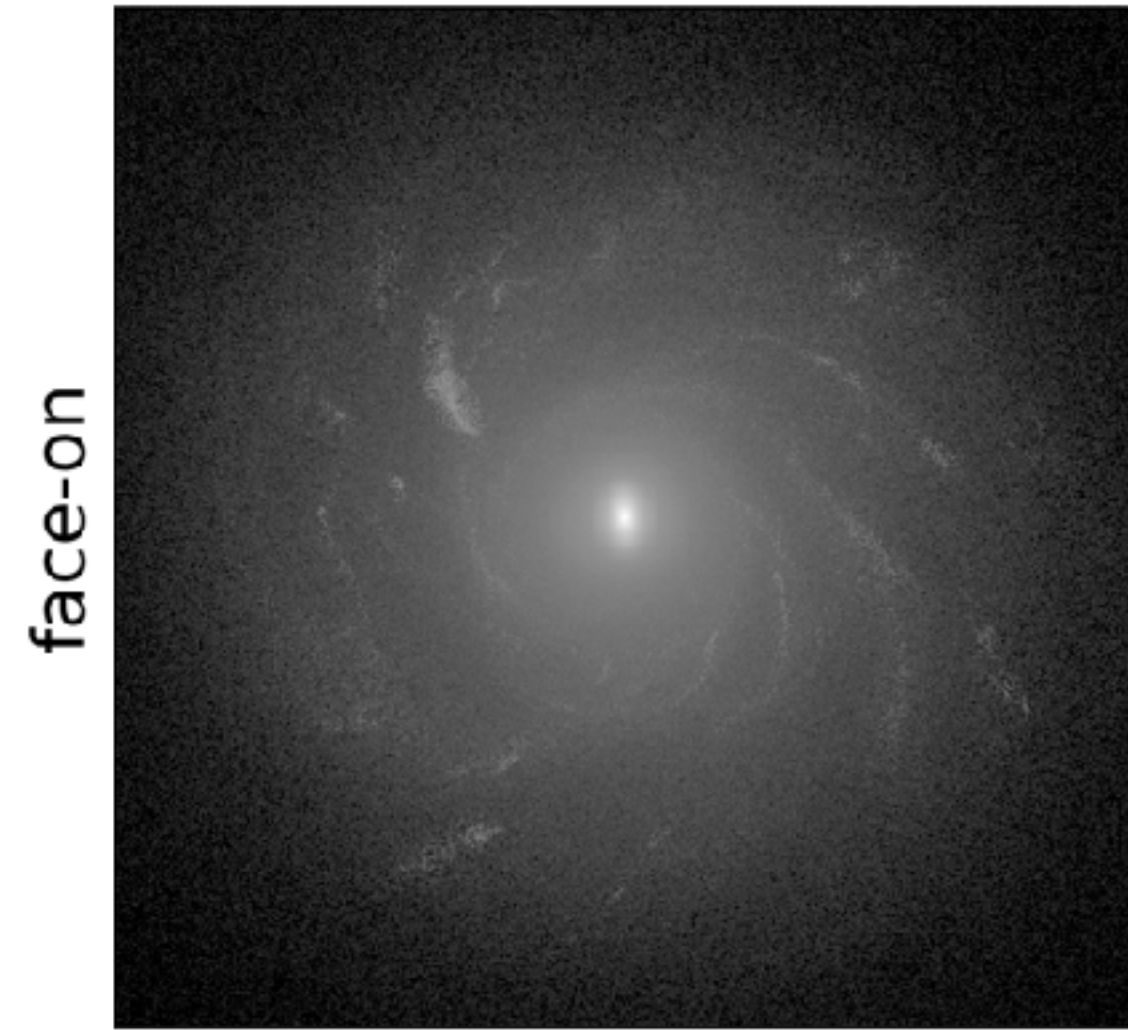
High-resolution Cluster Simulation

- Ramses Sciences at Yonsei 2018 -

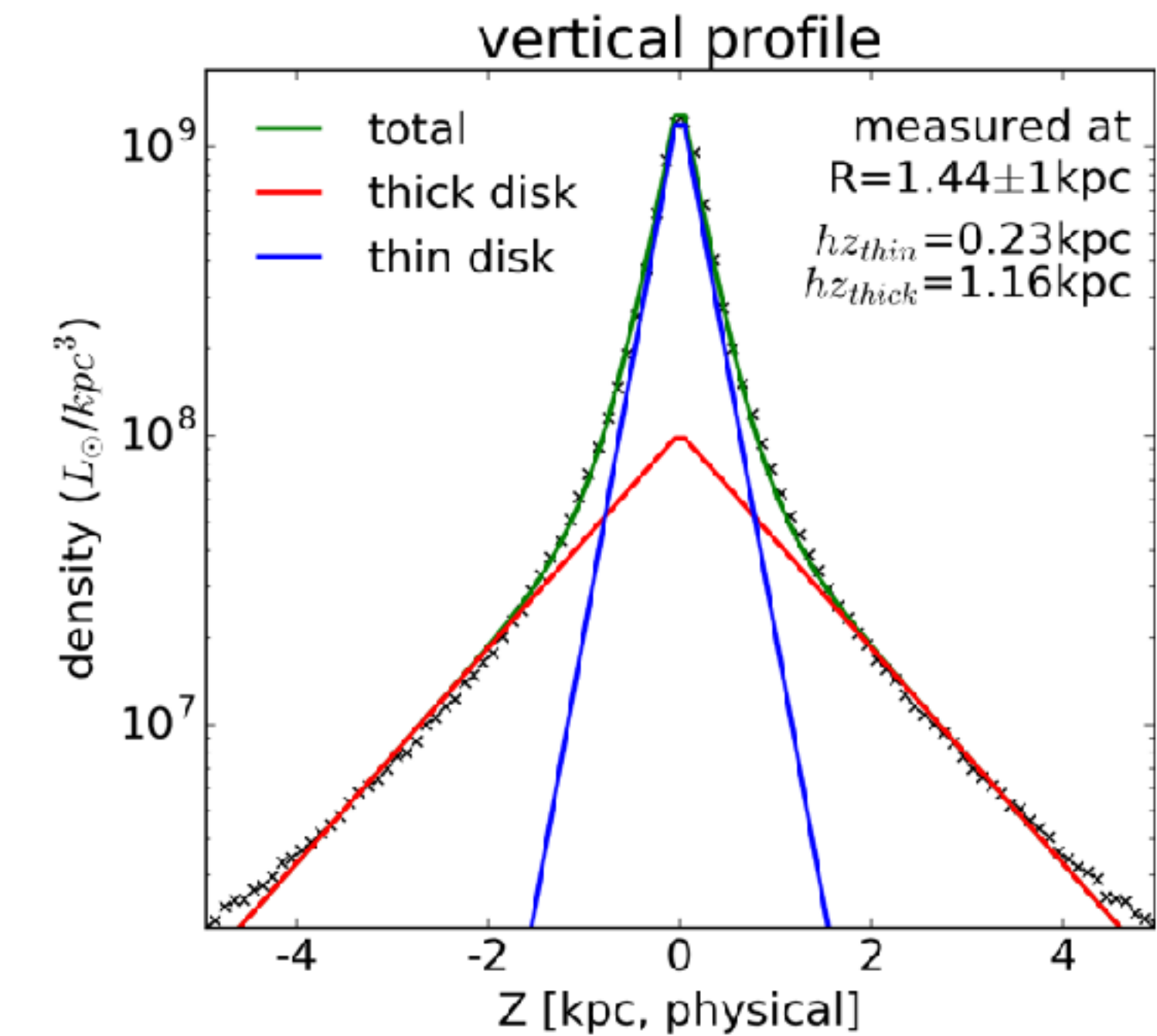
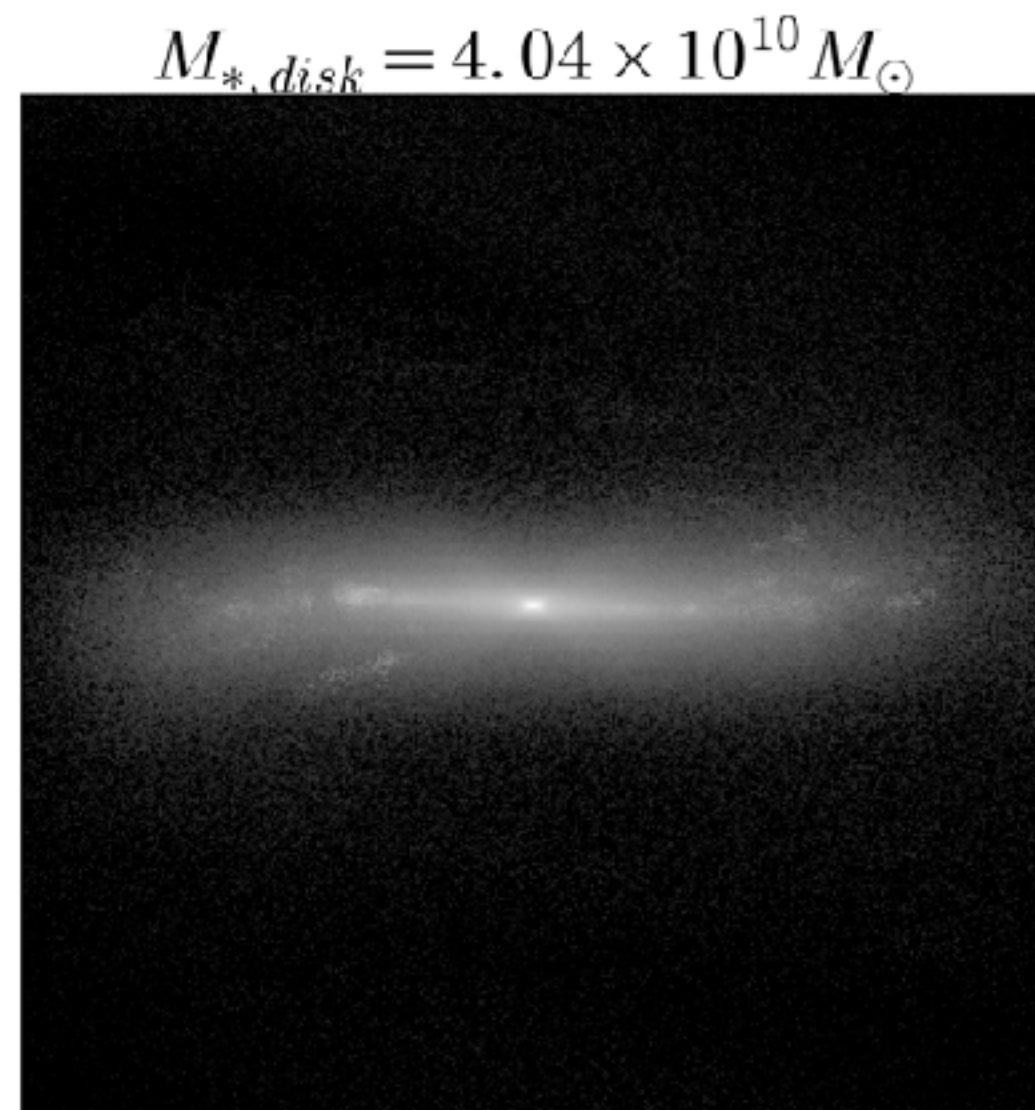
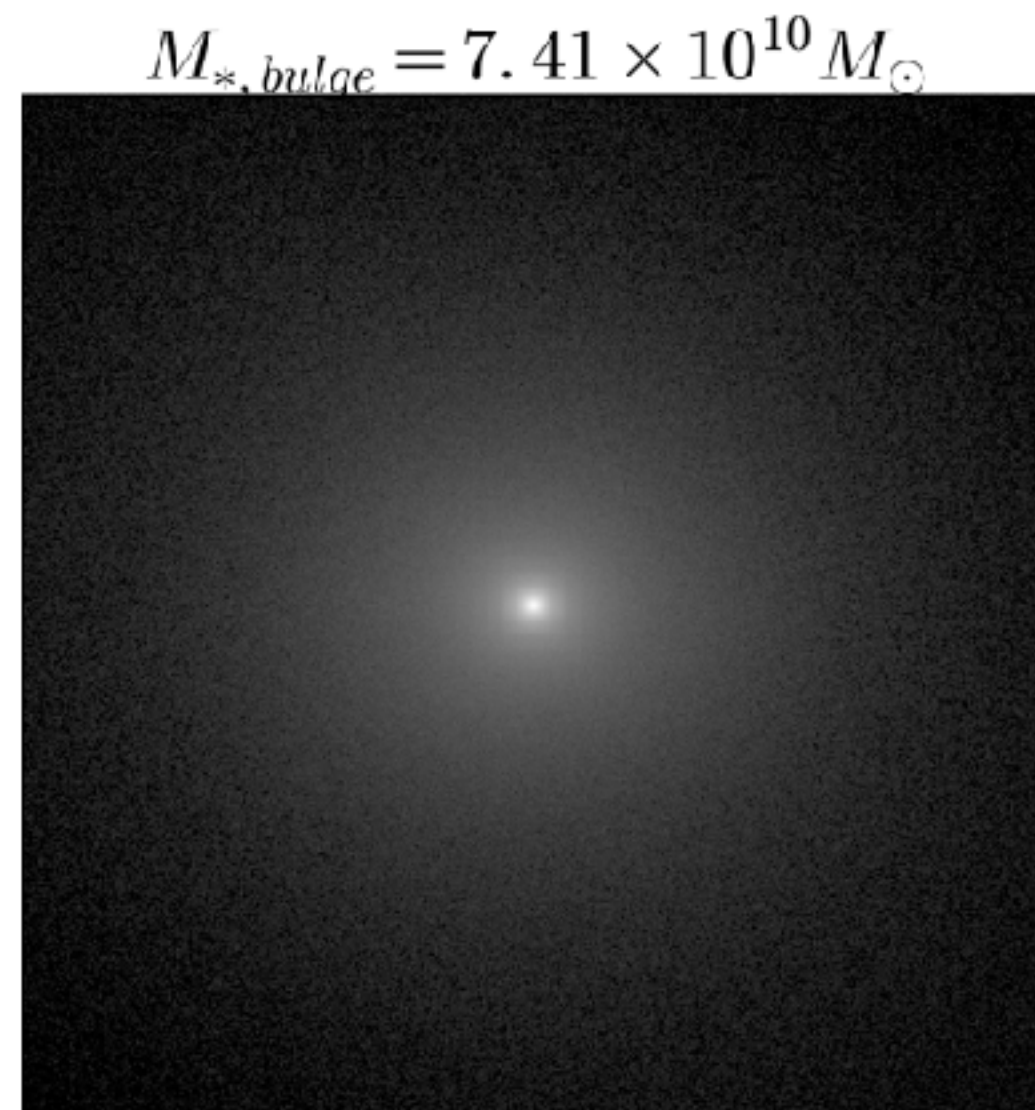
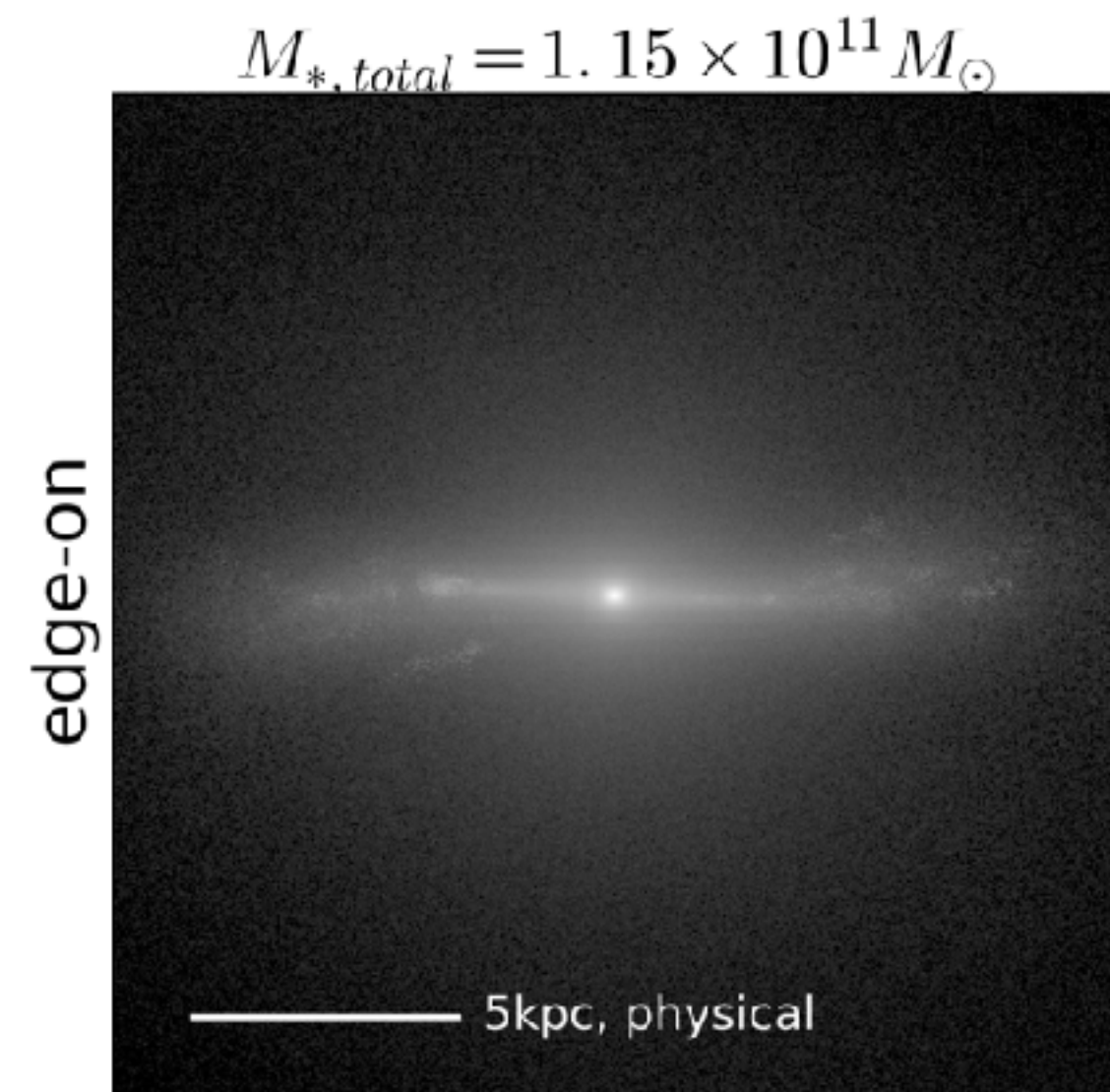
Sukyoung K. Yi

Disc settling



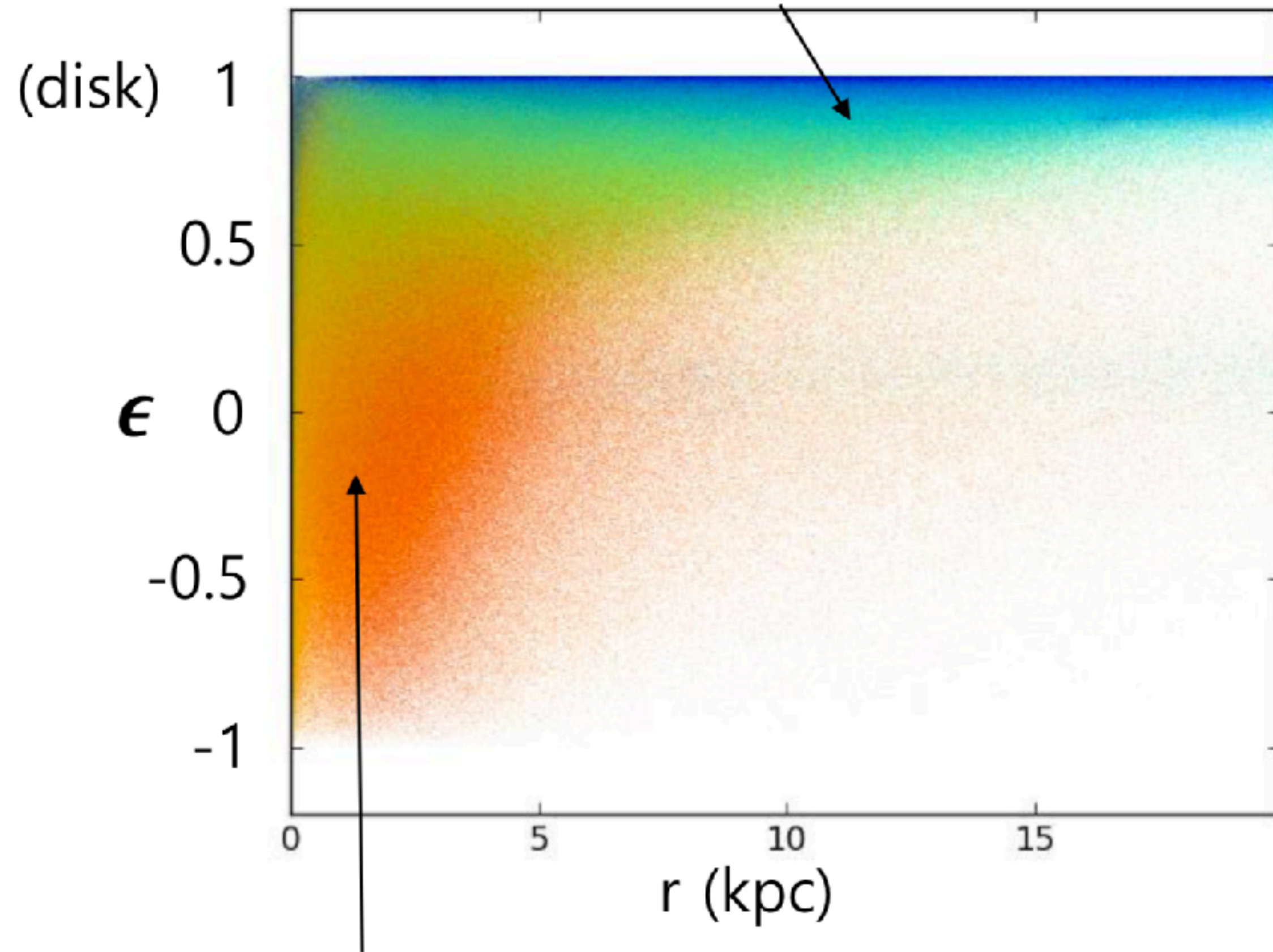


Disc resolved!



insitu SF

young stars $\epsilon \sim 1$ (disk)

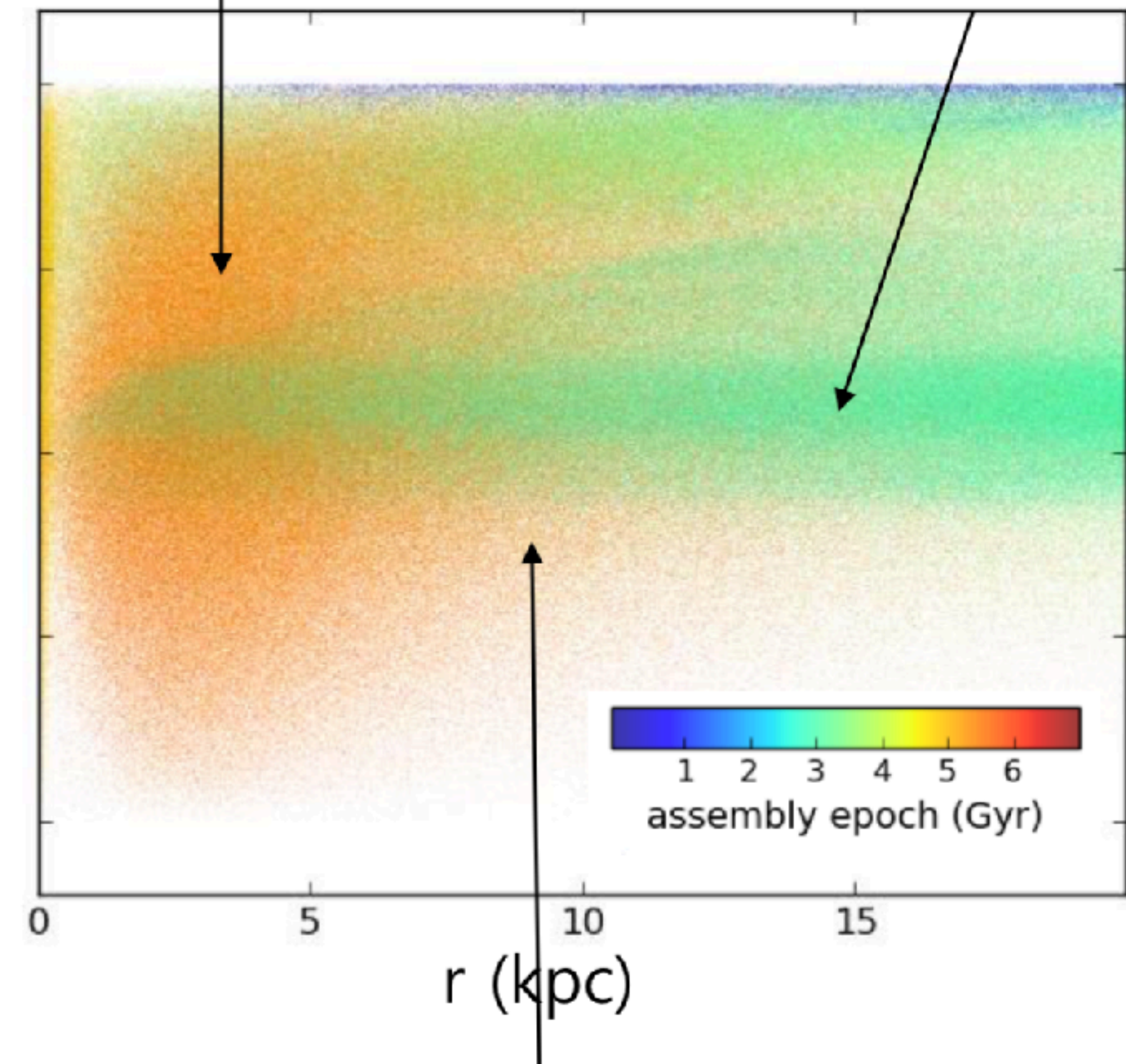


Old stars – disordered (inner regions)
“bulge”

accretion

merger
at $z \sim 3$

merger
at $z \sim 1.5$

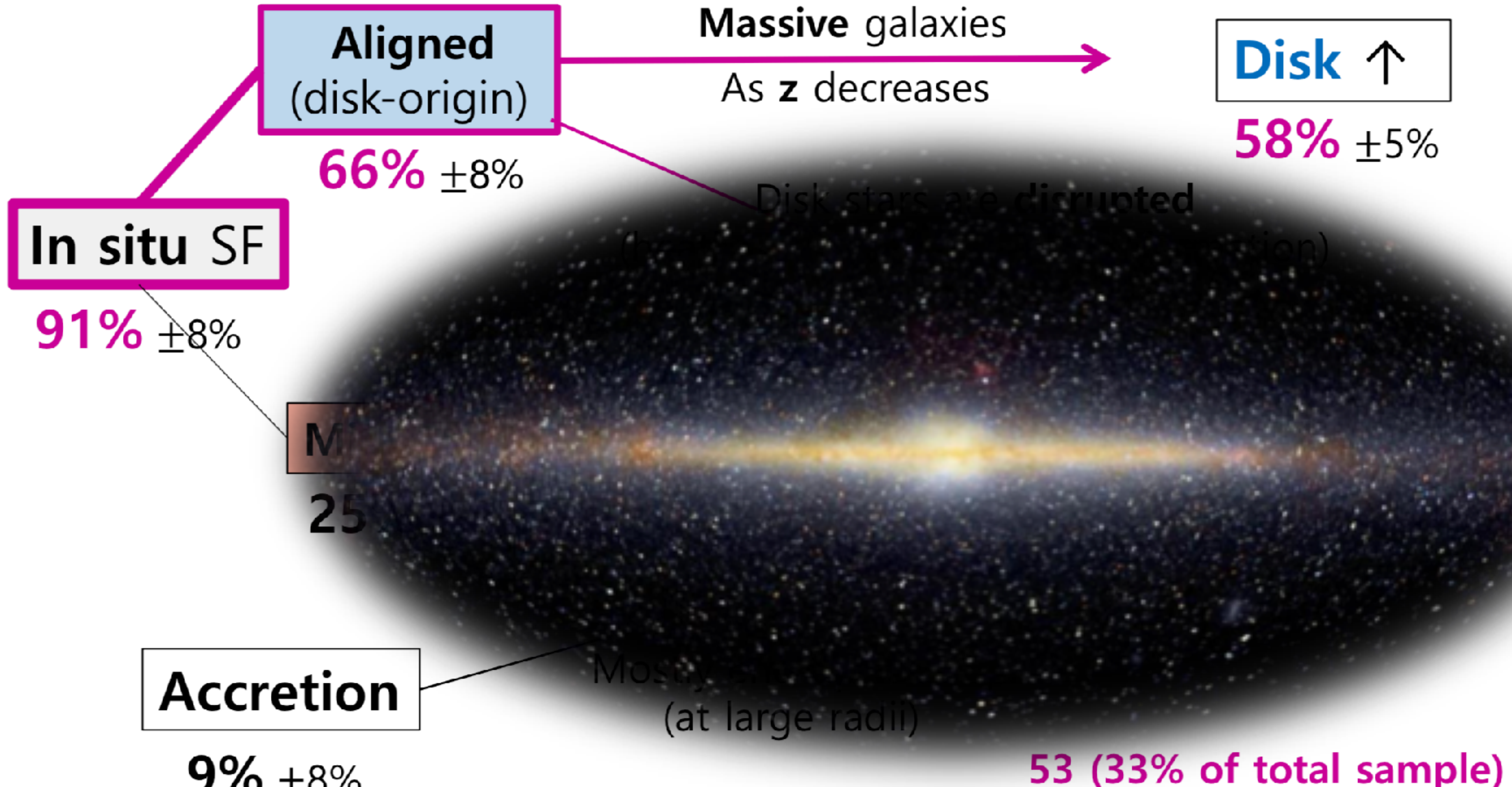


mostly disordered
recent accretion – larger radii

Origin of the structures of galaxies

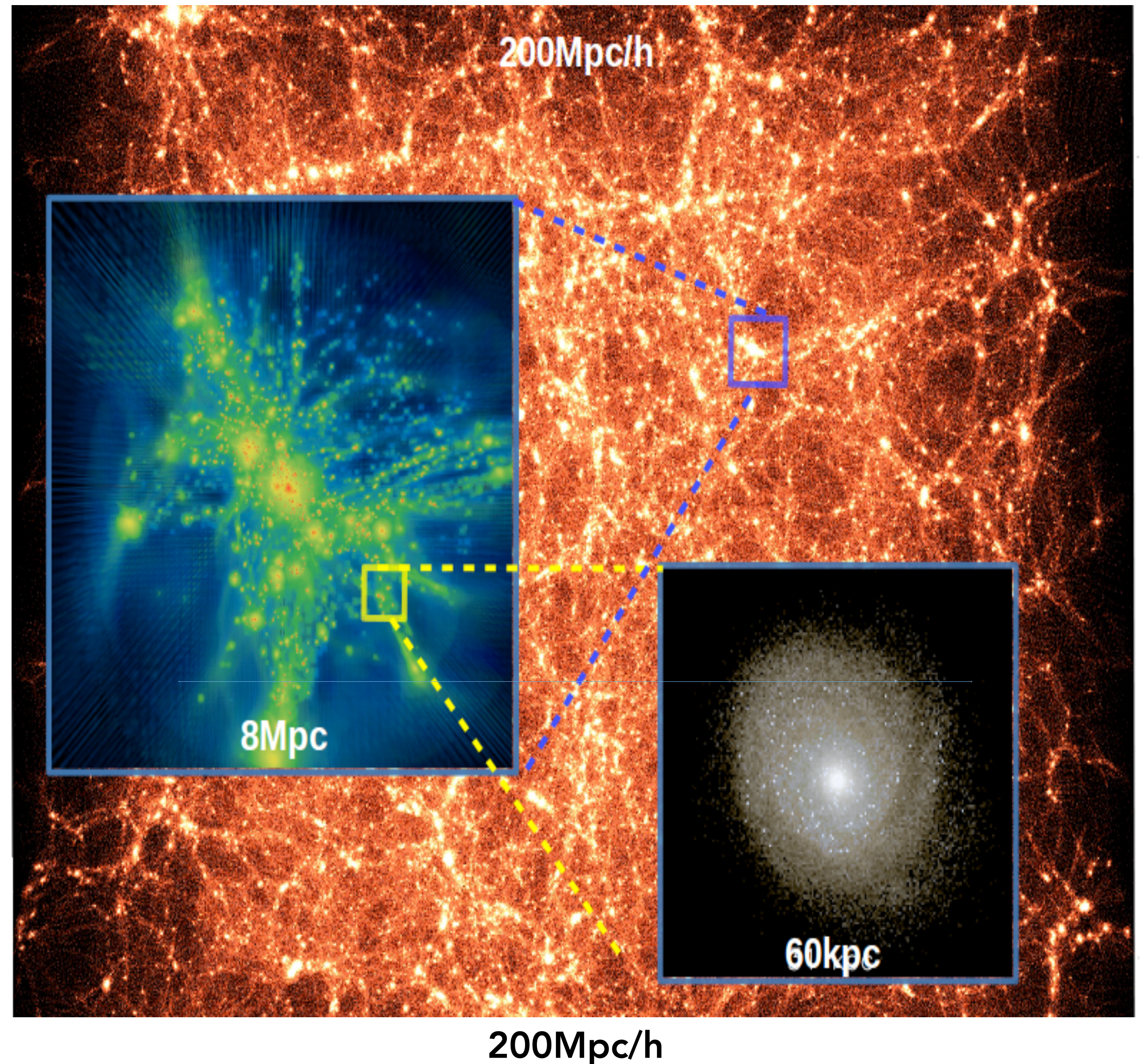
Minjung Park in prep

Measured at $z=0.7$



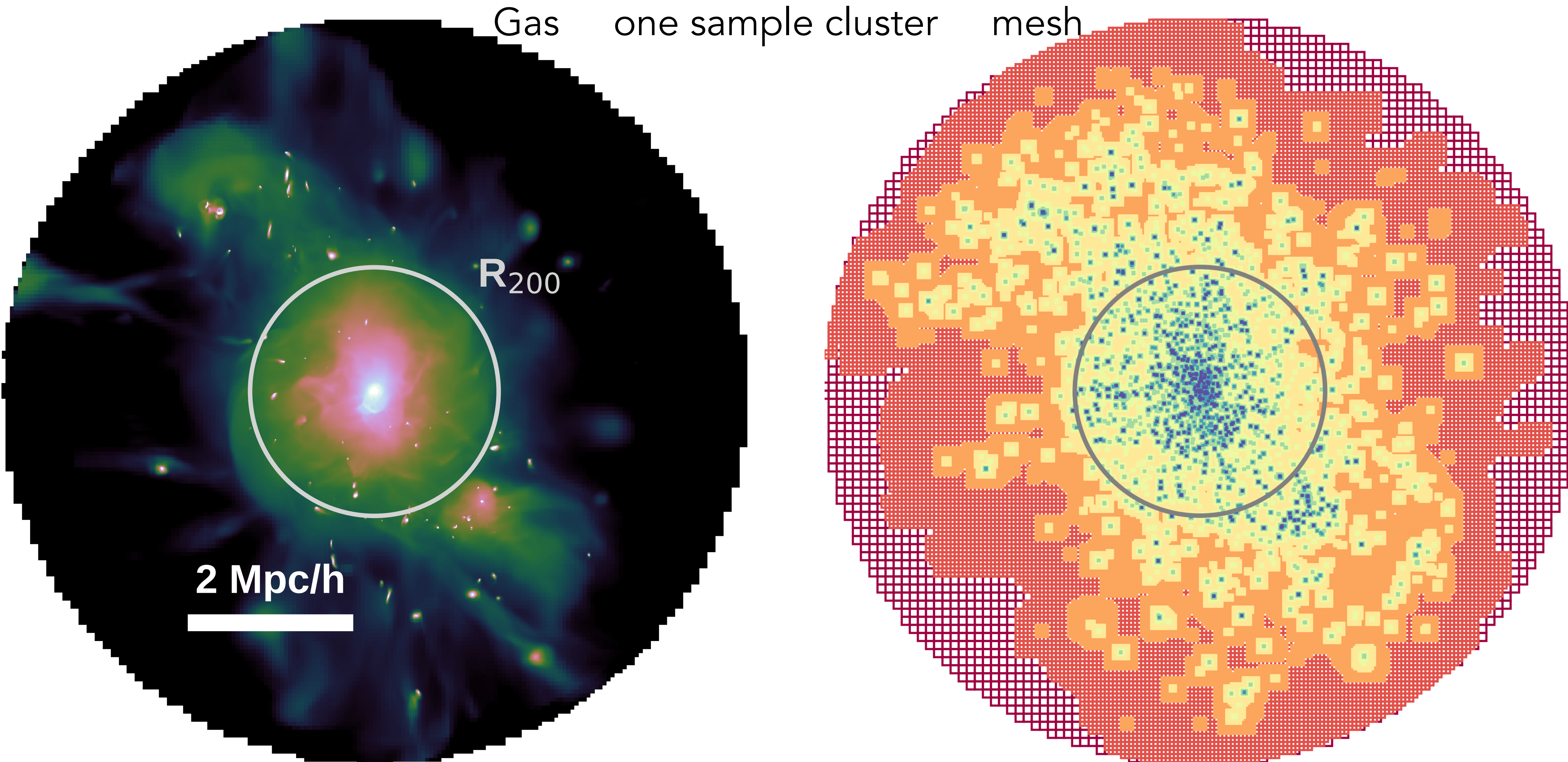
Yonsei Zoom-in Cluster Simulation (YZiCS)

- RAMSES (AMR, Teyssier 2002)
- KISTI 10M Chr
- Baryon recipes: SF, SN, AGN FB (HAGN: Dubois et al. 2014)
- 16 clusters in 200Mpc/h
- $13.5 < \log M/M_{\odot} < 15.0$
- $dm_{*}=4e5$
- $dm_{DM}=8e7$
- $dx = 0.76\text{kpc}/h$
- C.f. Bahe et al. 2013 (Gadget3, $dm_g \sim 1e6$, $dx \sim 1$, noAGN, $n \sim 30$ above 13.5)
- C.f. Barnes et al. 2017 (Gadget3, $dm_g \sim 2e6$, $dx \sim 1$, AGN, $n \sim 30$ above 14)

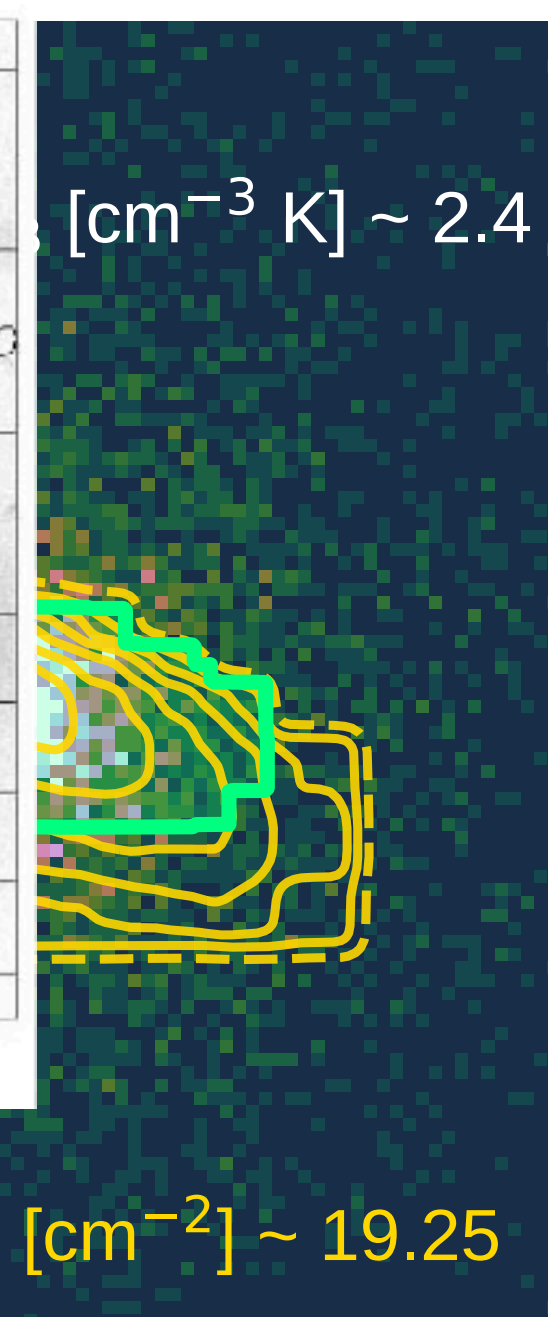
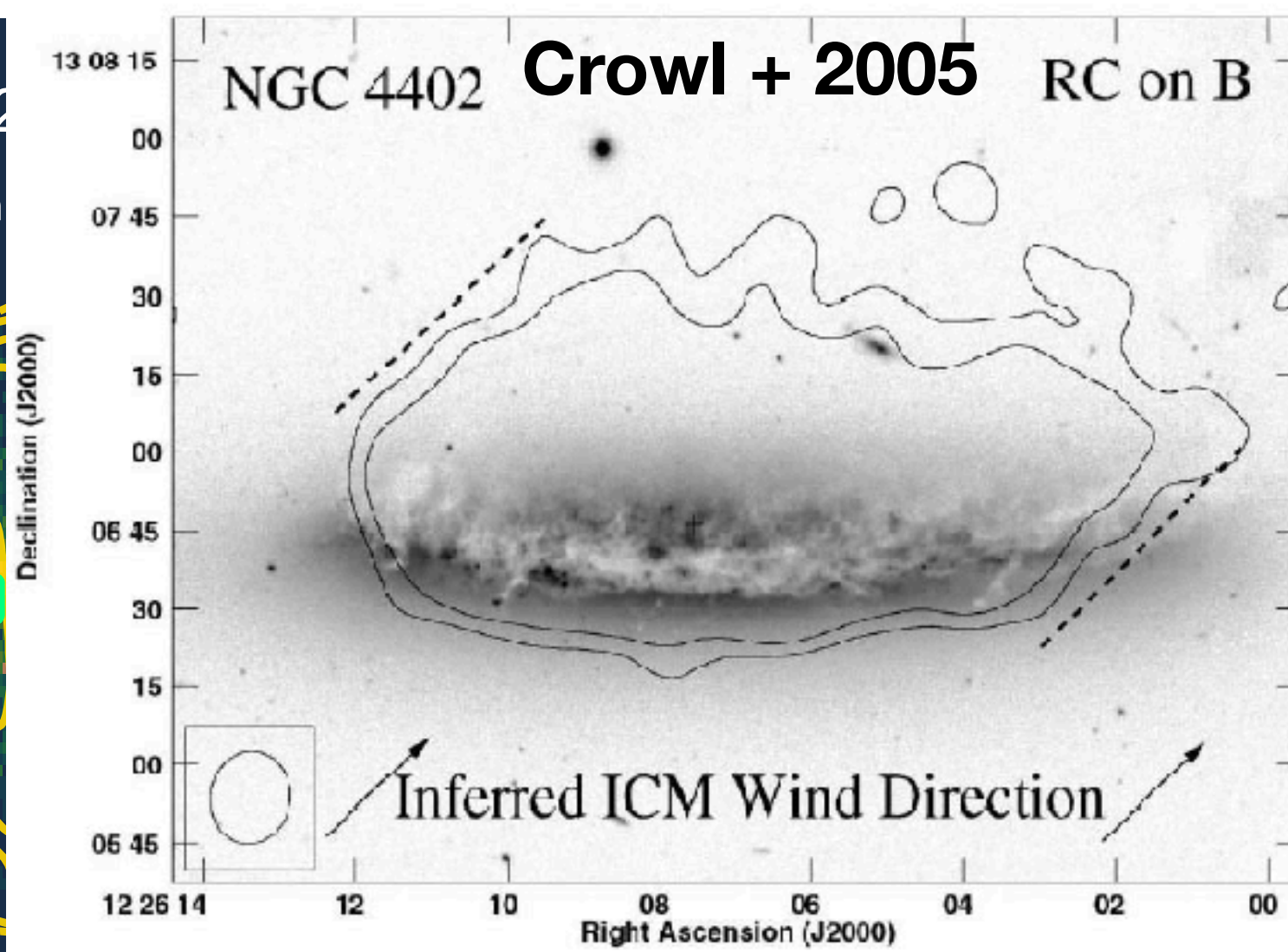
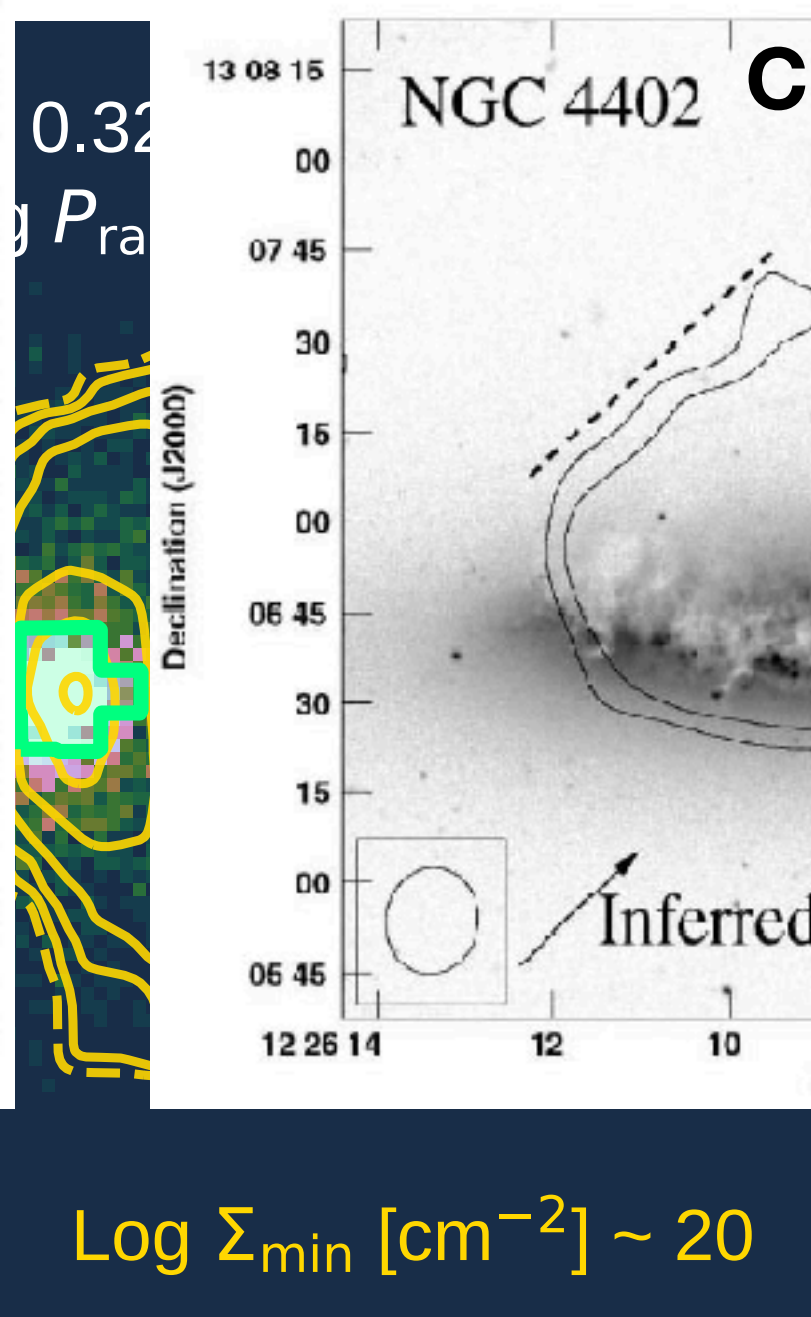
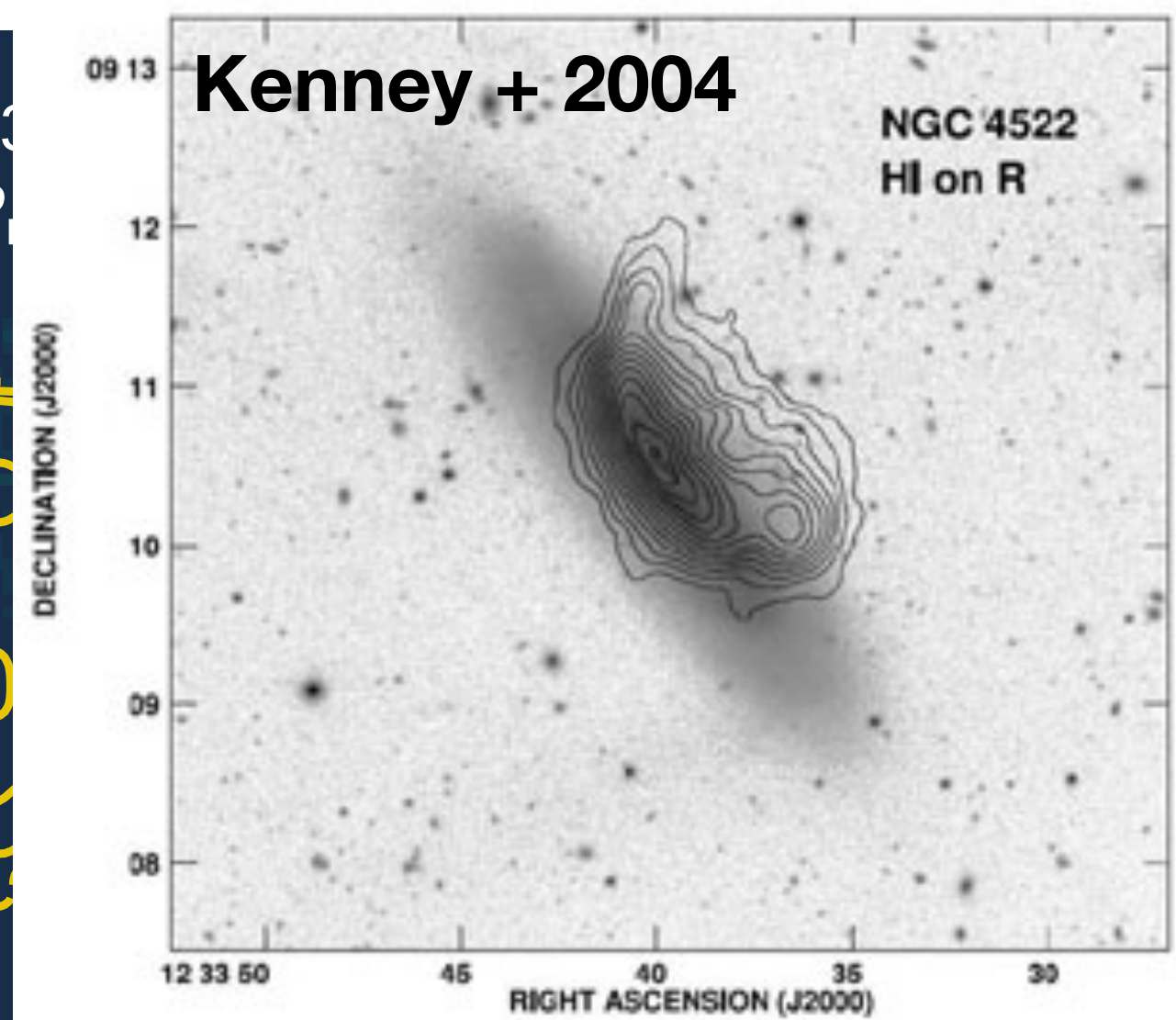
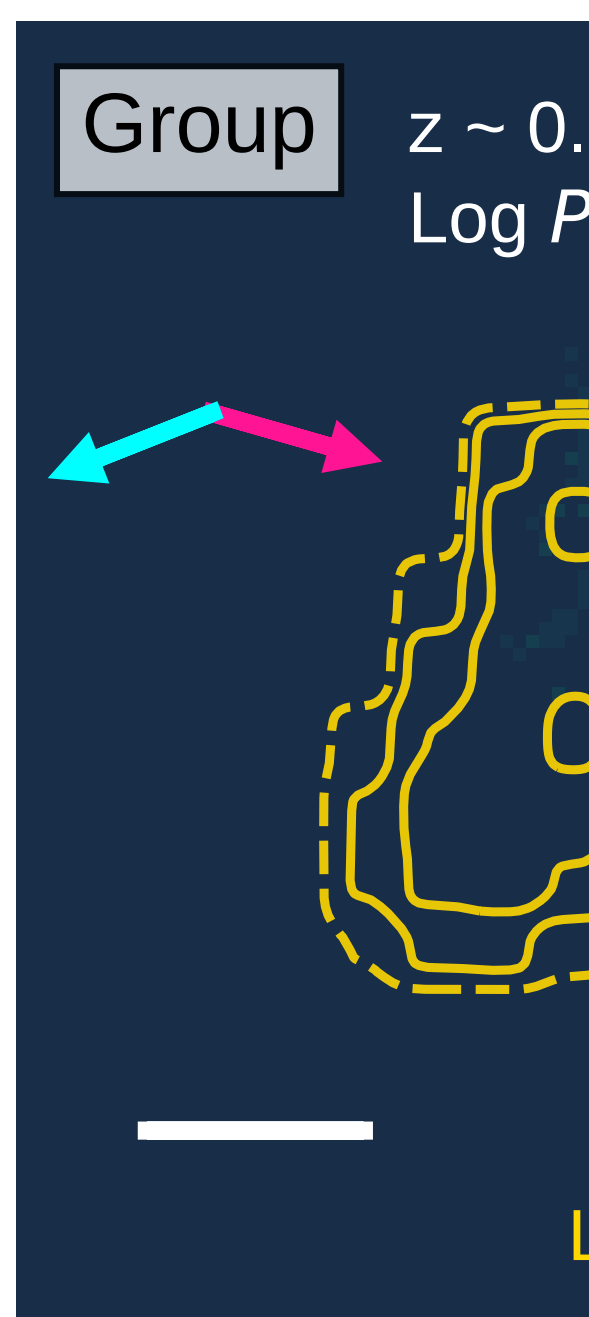
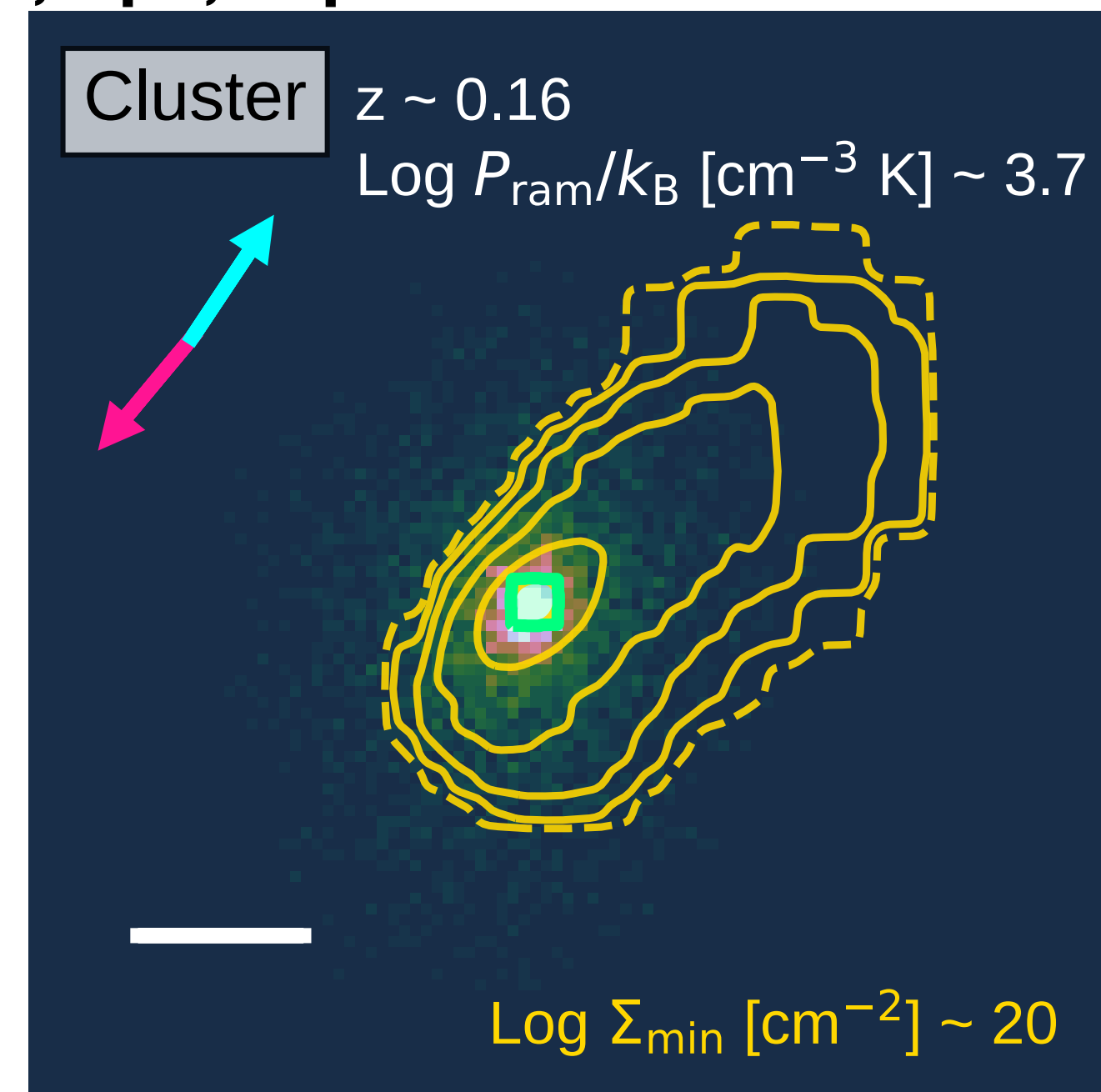
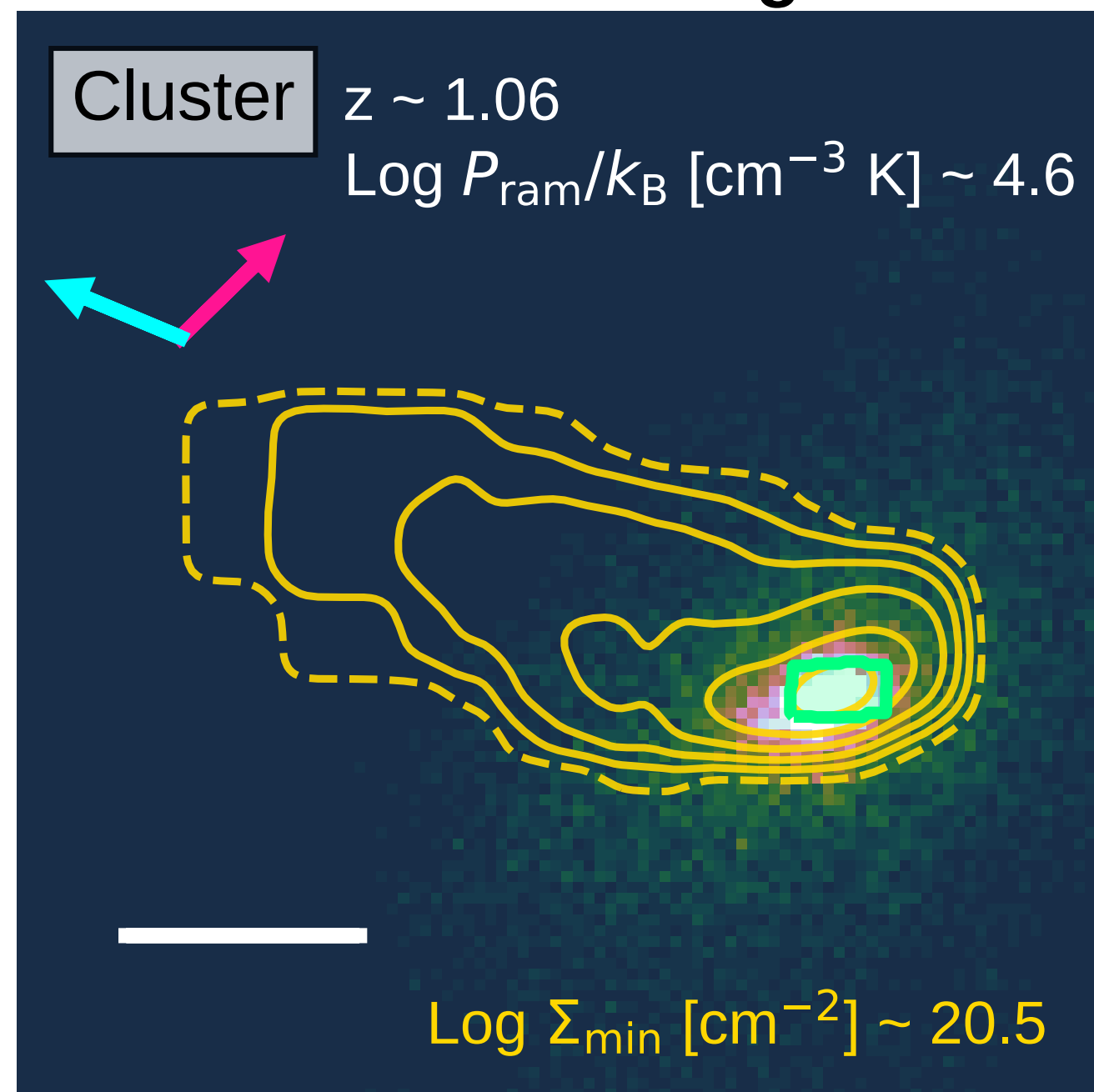
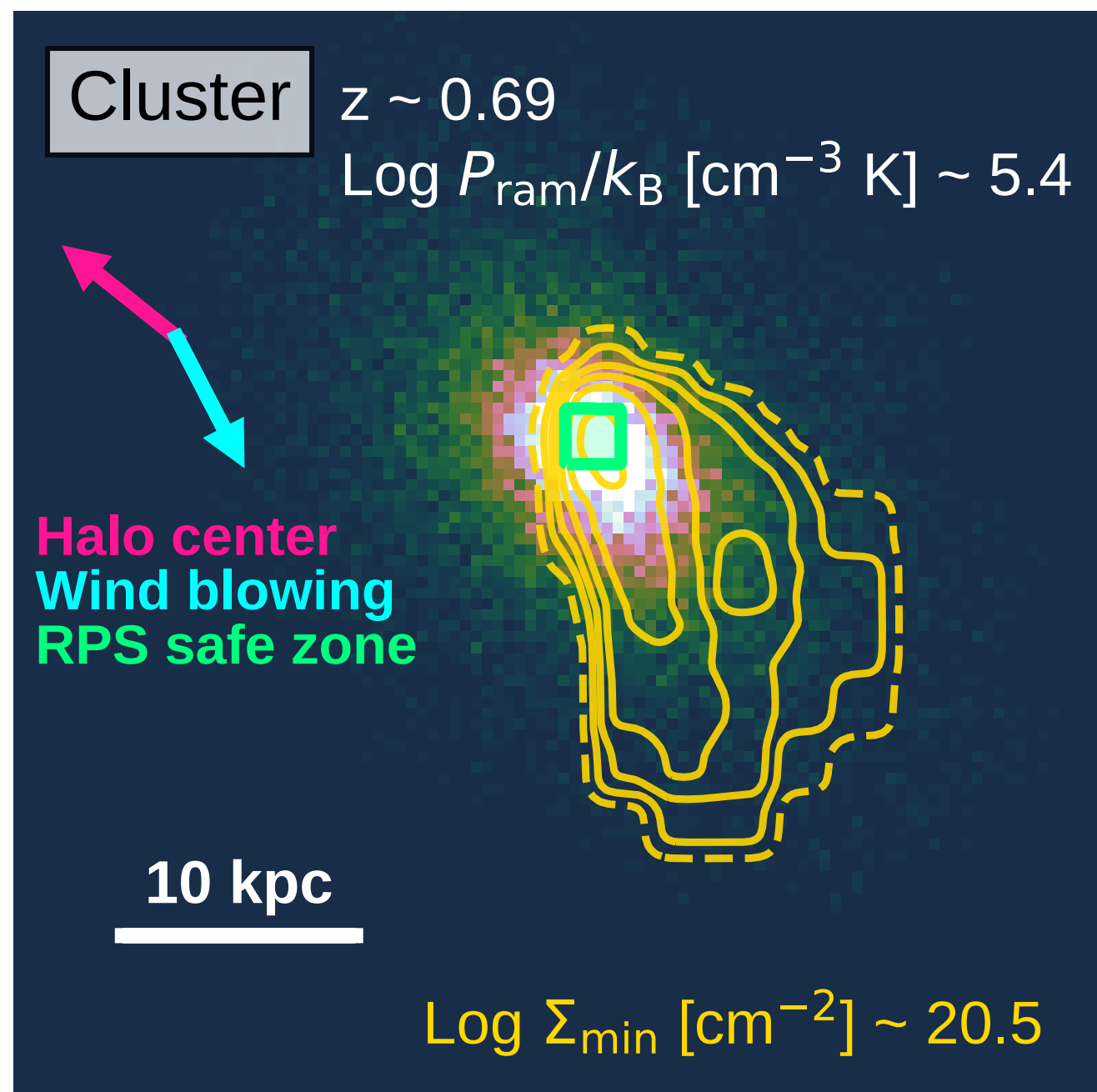


- Smith et al. 2016, ApJ, 833, 109 "Tidal stripping of DM and stars in galaxies"
- Choi & Yi 2017, ApJ, 837, 68 "Galaxy spin evolution in clusters"
- Rhee et al. 2017, ApJ, 843, 128 "Phase space analysis: Time since infall"
- Lee et al. 2018, ApJ, 864, 69 "Wobbling galaxy spin axes in dense environments"
- Jung et al. 2018, ApJ, in press "SF quenching of cluster galaxies":[astro-ph/1809.01684](https://arxiv.org/abs/1809.01684)
- Han et al. 2018, ApJ, in press "Preprocessing of dark haloes":[astro-ph/1809.02763](https://arxiv.org/abs/1809.02763)
- Rhee et al. 2018 ApJ, in prep "SF quenching of cluster galaxies using phase space analysis"

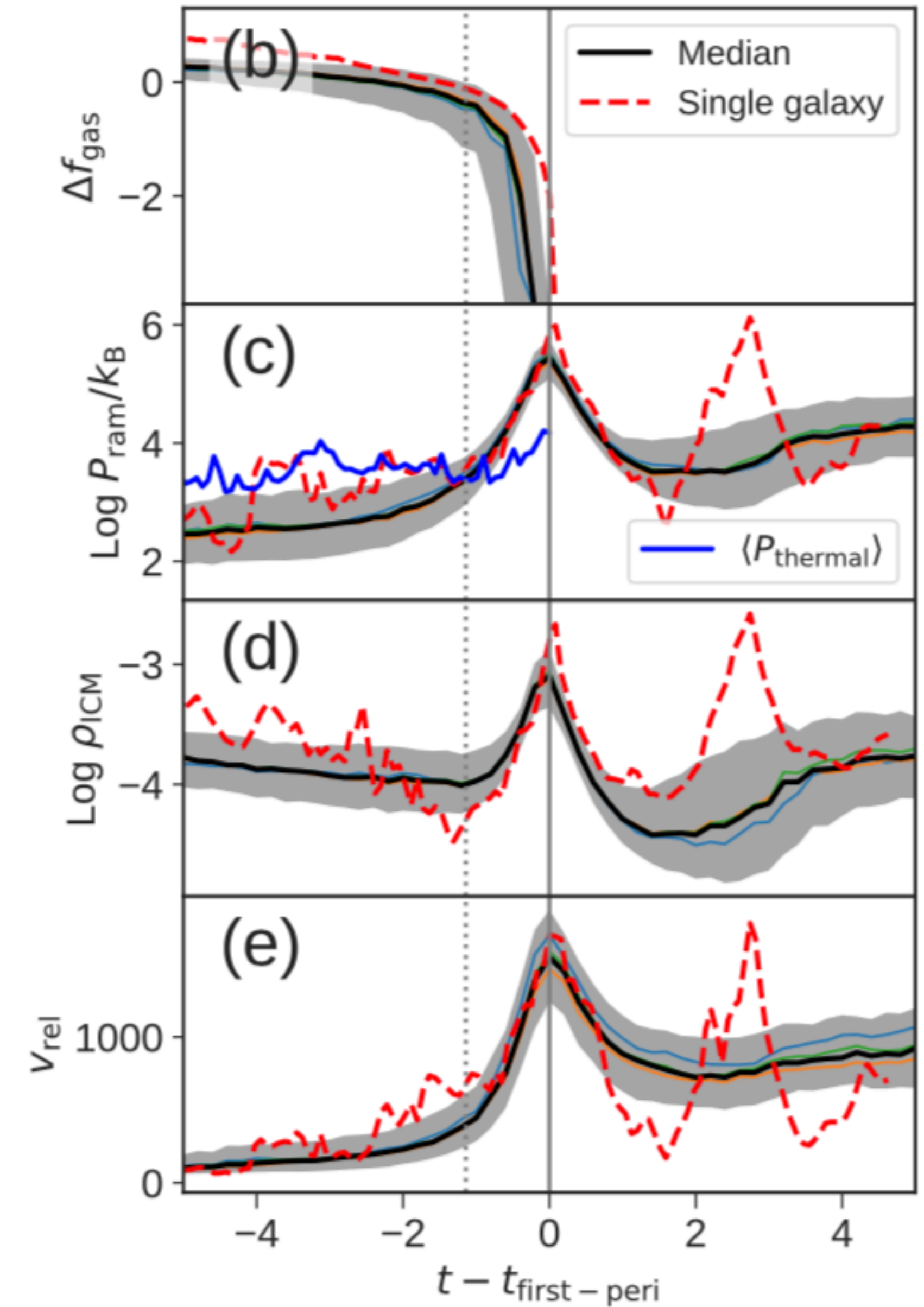
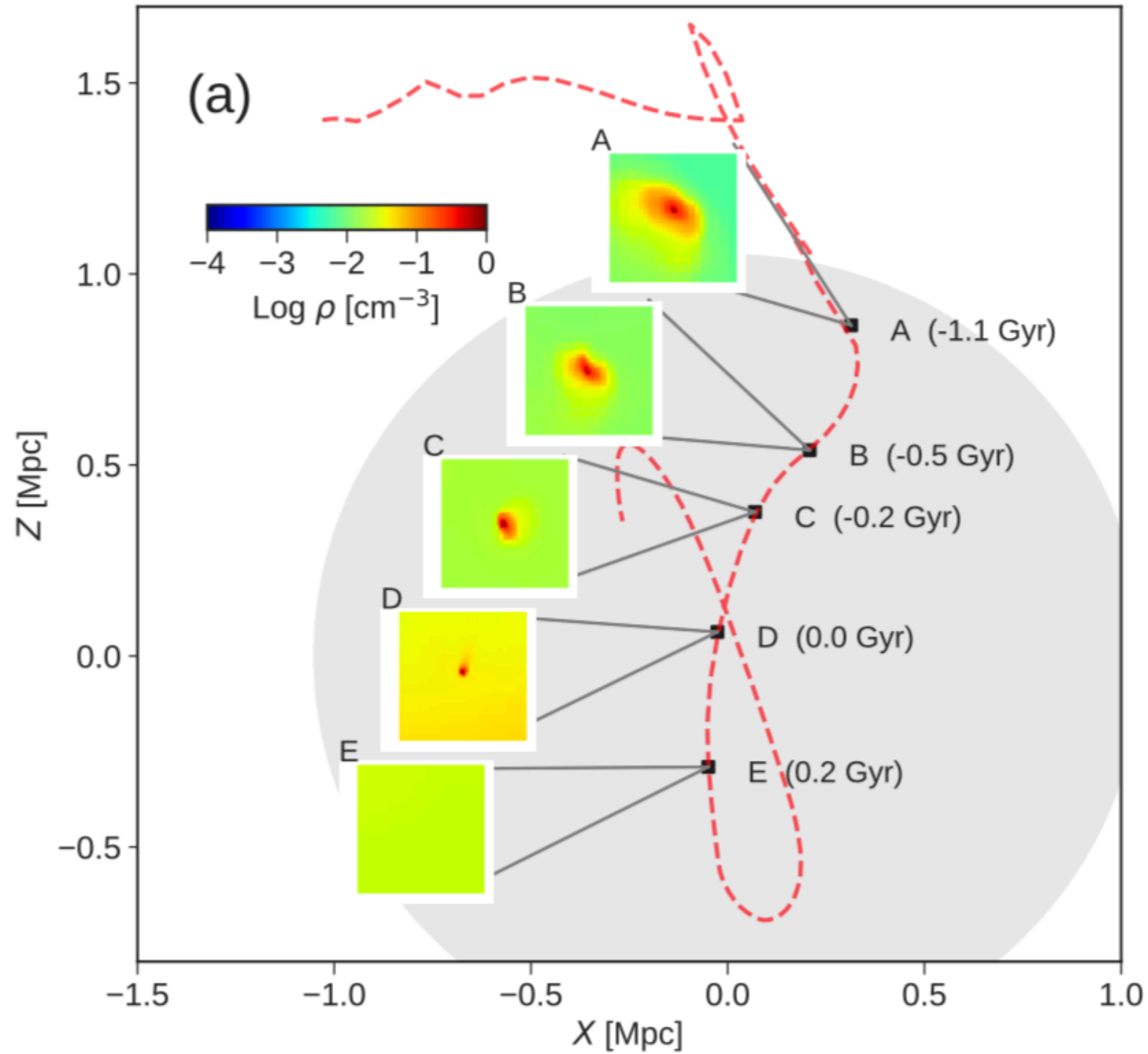
Resolution



Ram pressure stripping

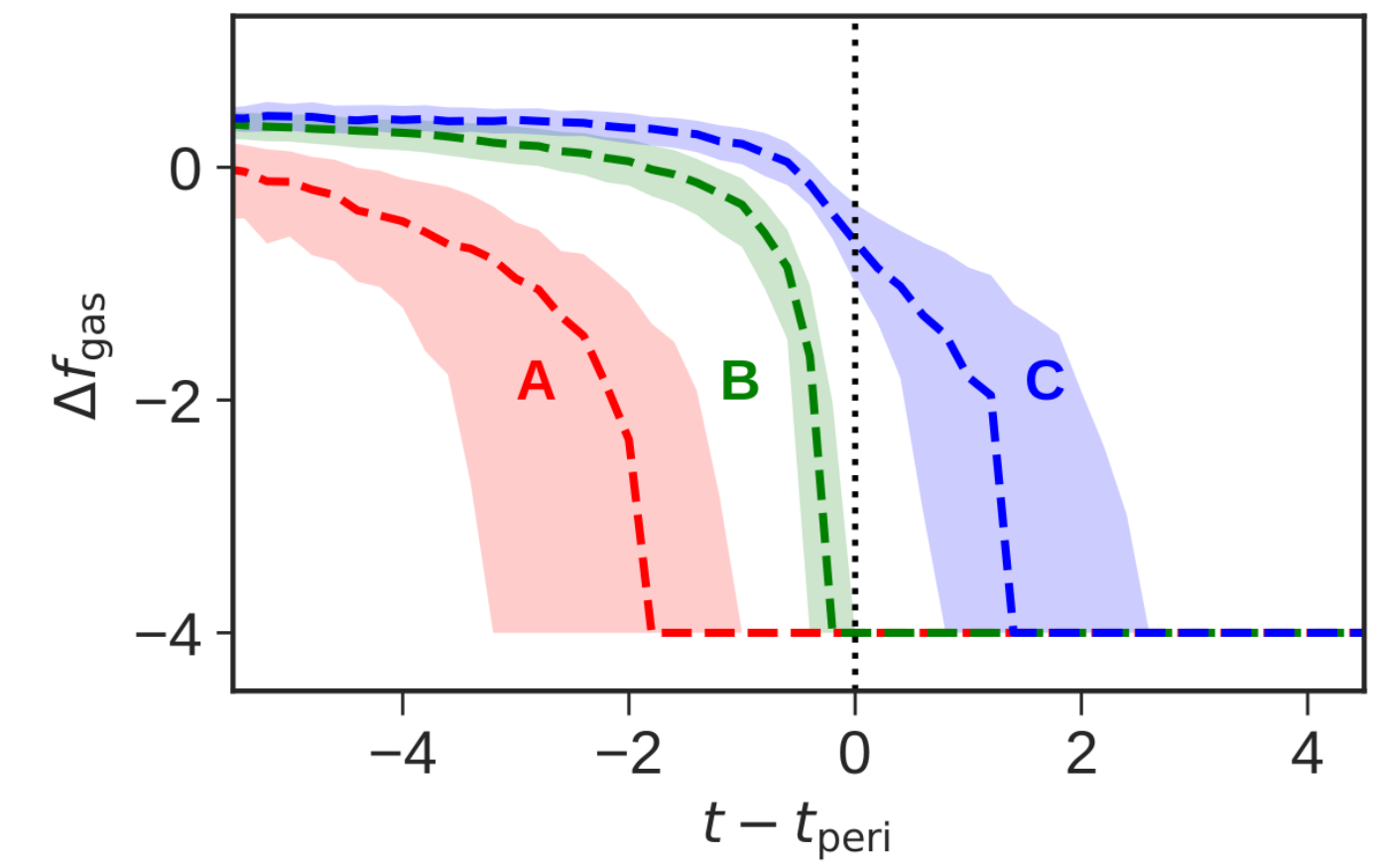
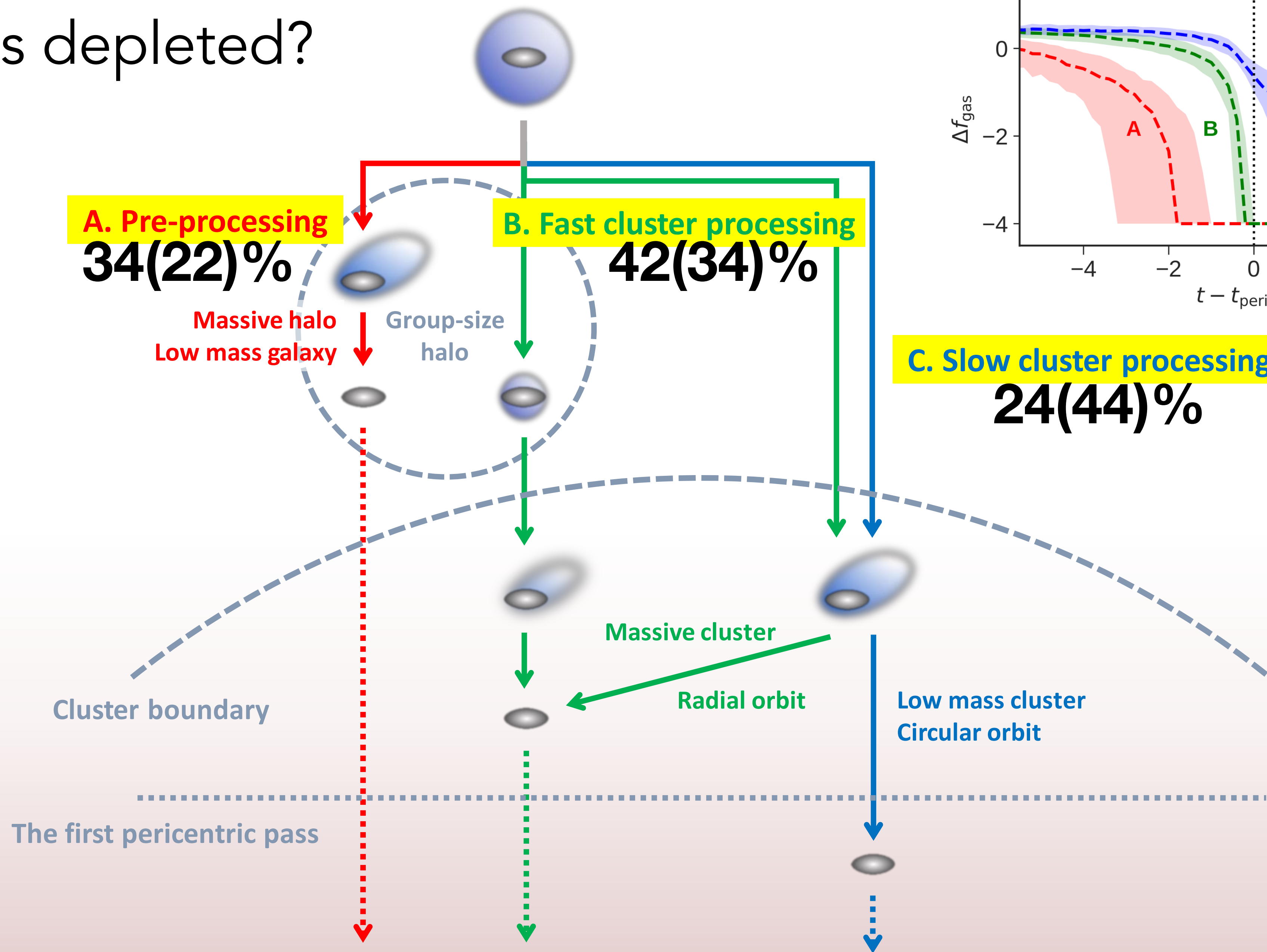


Ram pressure stripping



When/where do galaxies get gas depleted?

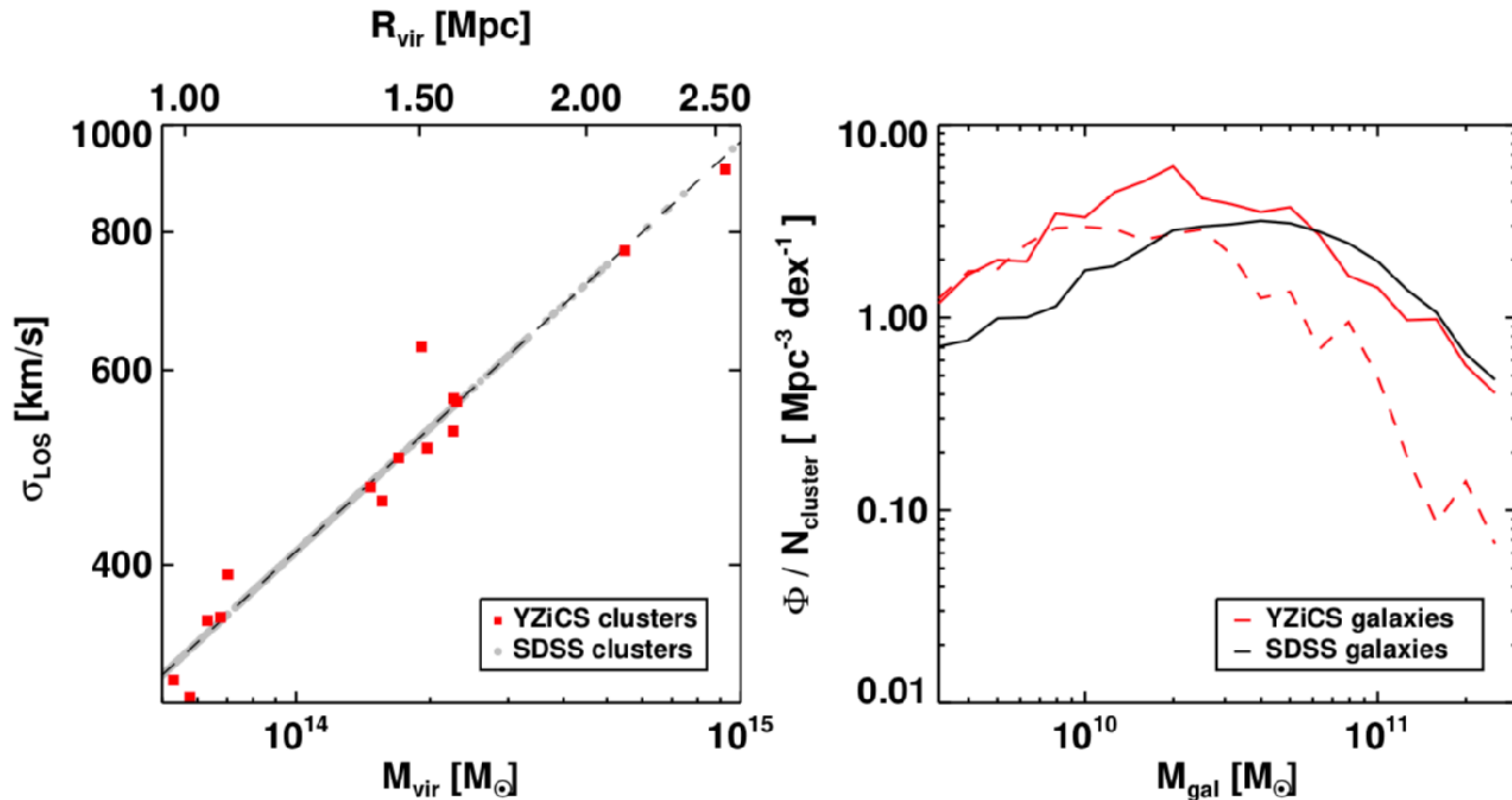
Jung et al. 2018, ApJ, in press: arXiv:1809.01684



C. Slow cluster processing
24(44)%

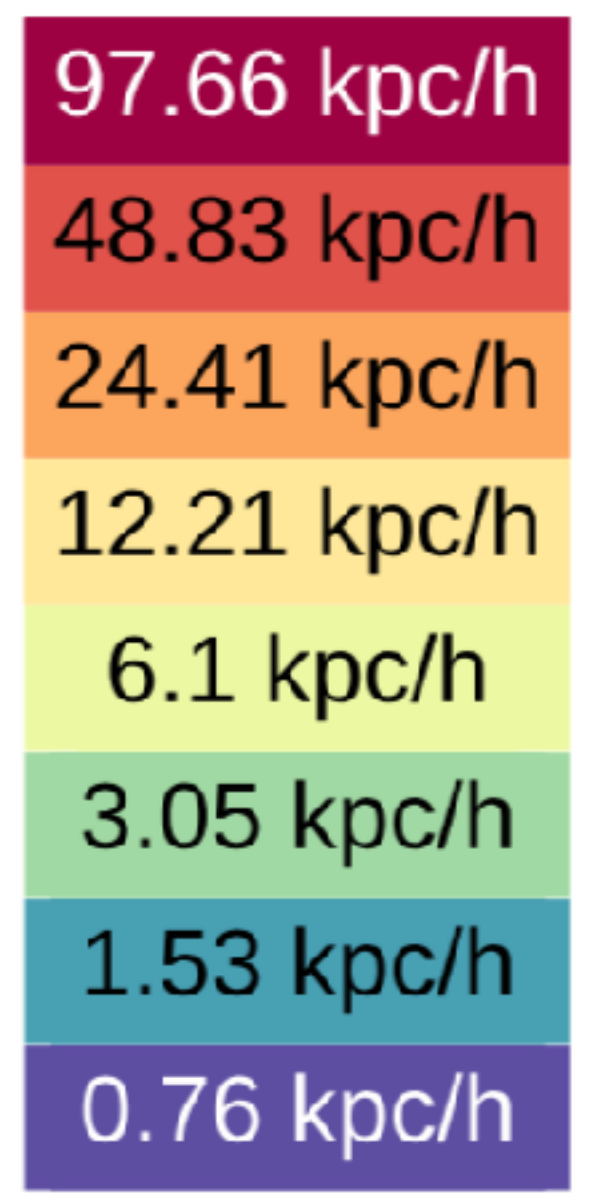
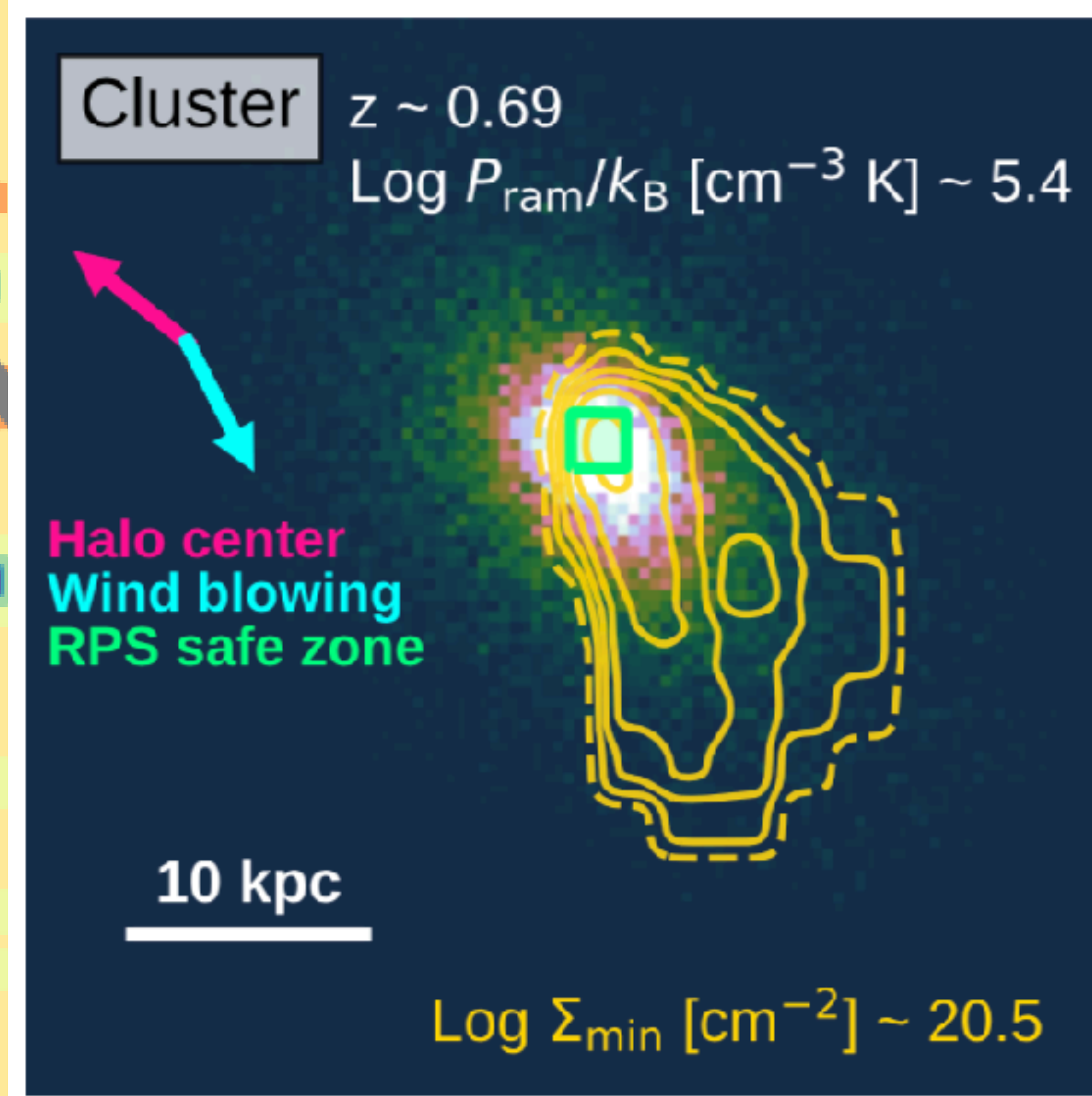
YZiCS clusters

YZiCS clusters are overly SF-quenched, ET, slow-rating: resolution, AGN, ?



C.f. C-Eagle (Barnes et al. 2017)
With AGN
galaxies are overly gas **rich**

YzICS resolution



Proposal for High-resolution cluster zoom-in simulation

- high res cluster simulation with NewHorizon specifications
- Tracer particles
- Alternative input ingredients (AGN, CR?)

KISTI Nurion (2018)

x85 more capability?



	GEN4	GEN5, Phi	GEN5, CPU
Processor	Xeon X5570 2.93GHz	KNL7250 1.4GHz	Xeon SKL6148 2.4GHz
Cores / Memory	8cores / 24GB	68cores / 96GB	40cores / 192GB
Node	3176	8305	132
Throughput	0.3PF (25,408 cores)	25.3PF (564,740 cores)	0.4PF (5,280 cores)
Total Memory	74TB	778TB	25TB
Interconnection	QDR(40Gbps)	OPI(100Gbps)	OPI(100Gbps)
Feature	SSE	AVX512 (8x)	AVX512 (8x)

"Super-challenge" with Nurion

- >5PF
- Xeon Phi processor, 68core/node ~ 3TF
- 5PF ~ 1700 nodes ~ 115,600 cores ~ 1700 x 96GB ~ 162TB ~ 7 x mem(NH)*

* mem(NH)=24TB

Ramses and Nurion

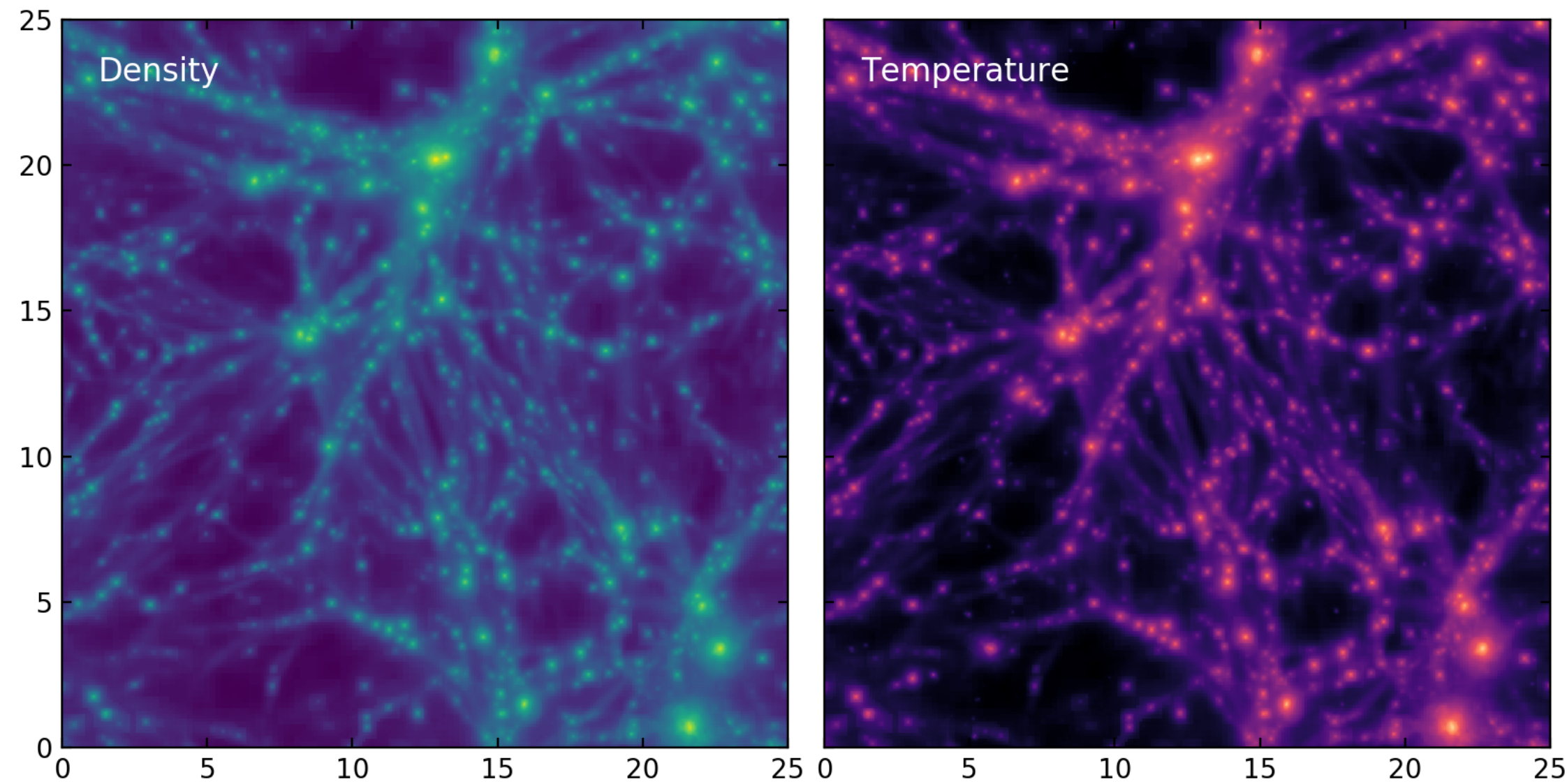
- “redundancy problem” of AMR
- “load balancing problem” of Ramses
- Ramses run on Nurion seems inefficient.
- mini Ramses?

Test run of mini-Ramses

- Performance test:
 - 25Mpc volume box, 128^3 DM particles, levelmin=7
 - 40 cores, Intel(R) Xeon(R) CPU 3.00GHz
 - Mini-RAMSES forces equipartition of AMR grids

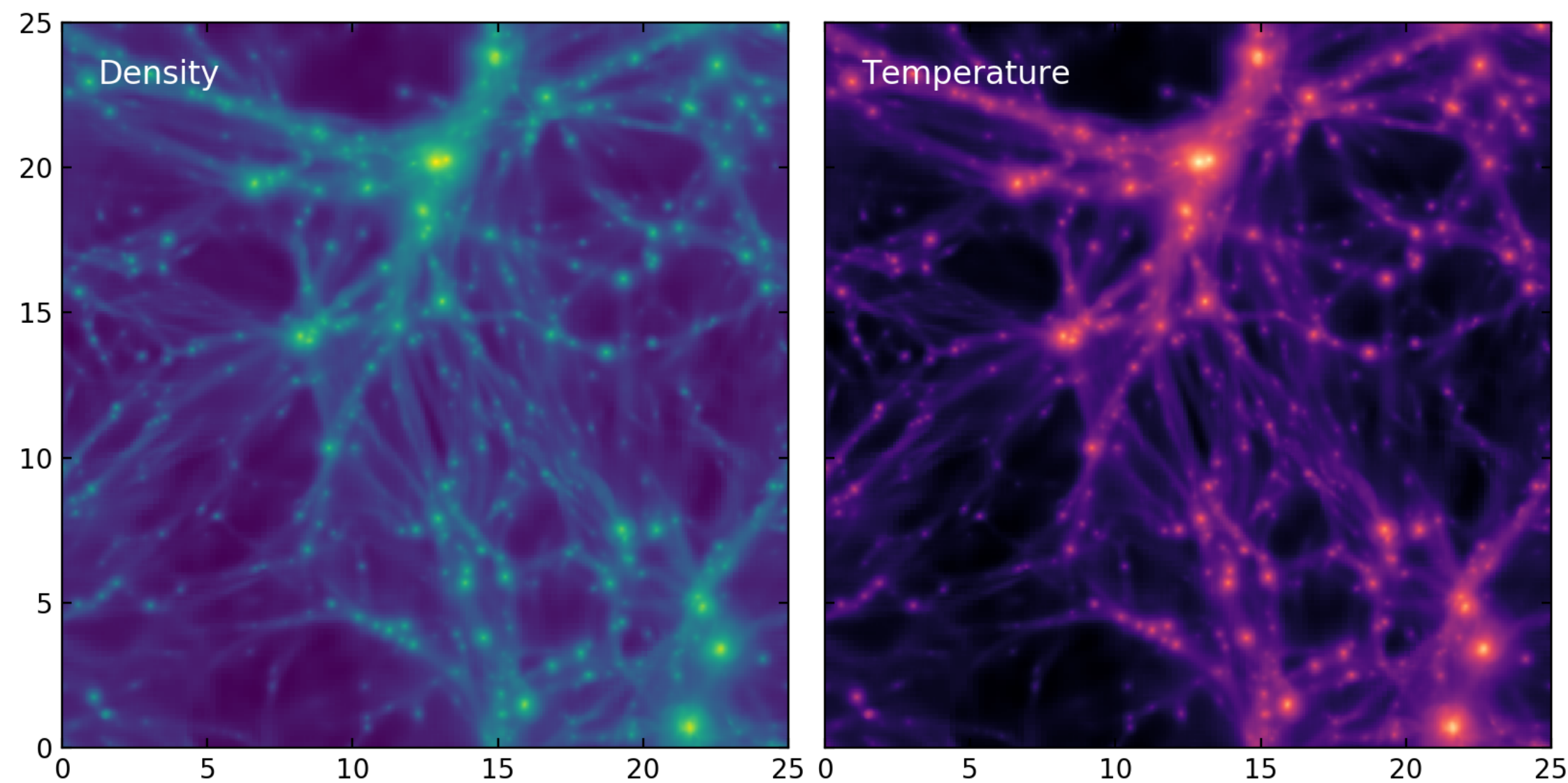
	RAMSES	Mini-RAMSES
Run time	104m 59s	184m 34s
Max. CPU grid	84942 (~0.12GB)	45038 (~0.06GB)

Ramses: DM and hydro only



(No stars, No cooling)

mini-Ramses:



X (Mpc/h)

- Detailed halo structure is different.
- Two codes may have different hydrodynamics scheme

Prospect on mini Ramses

- Is mini-Ramses superior in dealing with load balancing and redundancy?
- We are considering adding astrophysical prescriptions (cooling, SF, BH, FB, chemical enrichment, etc.) into mini-Ramses. Any duplicated effort?
- Is this doable?